



**Galaxy population properties of the massive X-ray  
luminous galaxy cluster XDCP J0044.0-2033 at  $z = 1.58$ .  
Red-sequence formation, massive galaxy assembly, and  
central star formation activity**

R. Fassbender, A. Nastasi, J. S. Santos, C. Lidman, M. Verdugo, Y. Koyama,  
P. Rosati, D. Pierini, N. Padilla, A. D. Romeo, et al.

► **To cite this version:**

R. Fassbender, A. Nastasi, J. S. Santos, C. Lidman, M. Verdugo, et al.. Galaxy population properties of the massive X-ray luminous galaxy cluster XDCP J0044.0-2033 at  $z = 1.58$ . Red-sequence formation, massive galaxy assembly, and central star formation activity. *Astronomy and Astrophysics - A&A*, 2014, 568, pp.A5. 10.1051/0004-6361/201423941 . cea-01271044

**HAL Id: cea-01271044**

**<https://hal-cea.archives-ouvertes.fr/cea-01271044>**

Submitted on 8 Feb 2016

**HAL** is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

# Galaxy population properties of the massive X-ray luminous galaxy cluster XDCP J0044.0-2033 at $z = 1.58^{*,**}$

## Red-sequence formation, massive galaxy assembly, and central star formation activity

R. Fassbender<sup>1,2</sup>, A. Nastasi<sup>3,2</sup>, J. S. Santos<sup>4,5</sup>, C. Lidman<sup>6</sup>, M. Verdugo<sup>7</sup>, Y. Koyama<sup>8,9</sup>, P. Rosati<sup>10,11</sup>,  
 D. Pierini<sup>2</sup>, N. Padilla<sup>12</sup>, A. D. Romeo<sup>13</sup>, N. Menci<sup>1</sup>, A. Bongiorno<sup>1</sup>, M. Castellano<sup>1</sup>, P. Cerulo<sup>15</sup>,  
 A. Fontana<sup>1</sup>, A. Galametz<sup>1</sup>, A. Grazian<sup>1</sup>, A. Lamastra<sup>1</sup>, L. Pentericci<sup>1</sup>, V. Sommariva<sup>1</sup>, V. Strazzullo<sup>14</sup>,  
 R. Šuhada<sup>16</sup>, and P. Tozzi<sup>4</sup>

(Affiliations can be found after the references)

Received 3 April 2014 / Accepted 5 June 2014

### ABSTRACT

**Context.** Recent observational progress has enabled the detection of galaxy clusters and groups out to very high redshifts and for the first time allows detailed studies of galaxy population properties in these densest environments in what was formerly known as the “redshift desert” at  $z > 1.5$ .

**Aims.** We aim to investigate various galaxy population properties of the massive X-ray luminous galaxy cluster XDCP J0044.0-2033 at  $z = 1.58$ , which constitutes the most extreme currently known matter-density peak at this redshift.

**Methods.** We analyzed deep VLT/HAWK-I near-infrared data with an image quality of  $0.5''$  and limiting Vega magnitudes (50% completeness) of 24.2 in  $J$ - and 22.8 in the  $K_s$  band, complemented by similarly deep Subaru imaging in  $i$  and  $V$ , *Spitzer* observations at  $4.5 \mu\text{m}$ , and new spectroscopic observations with VLT/FORS 2.

**Results.** We detect a cluster-associated excess population of about 90 galaxies, most of them located within the inner  $30''$  (250 kpc) of the X-ray centroid, which follows a centrally peaked, compact NFW galaxy surface-density profile with a concentration of  $c_{200} \approx 10$ . Based on the *Spitzer*  $4.5 \mu\text{m}$  imaging data, we measure a total enclosed stellar mass of  $M_{*500} \approx (6.3 \pm 1.6) \times 10^{12} M_{\odot}$  and a resulting stellar mass fraction of  $f_{*,500} = M_{*,500}/M_{500} = (3.3 \pm 1.4)\%$ , consistent with local values. The total  $J$ - and  $K_s$ -band galaxy luminosity functions of the core region yield characteristic magnitudes  $J^*$  and  $K_s^*$  consistent with expectations from simple  $z_f = 3$  burst models. However, a detailed look at the morphologies and color distributions of the spectroscopically confirmed members reveals that the most massive galaxies are undergoing a very active mass-assembly epoch through merging processes. Consequently, the bright end of the cluster red sequence is not in place, while a red-locus population is present at intermediate magnitudes  $[K_s^*, K_s^* + 1.6]$ , which is then sharply truncated at magnitudes fainter than  $K_s^* + 1.6$ . The dominant cluster-core population comprises post-quenched galaxies transitioning toward the red sequence at intermediate magnitudes, while additionally a significant blue-cloud population of faint star-forming galaxies is present even in the densest central regions. Based on a color–color selection performed to separate different cluster galaxy types, we find that the blue star-forming population is concentrated in clumpy structures and dominates in particular at and beyond the  $R_{500}$  radius. On the other hand, the fraction of post-starburst galaxies steadily increases toward the center, while the red-locus population and red-sequence transition galaxies seem to reach their peak fractions already at intermediate cluster-centric radii of about  $r \sim 200$  kpc.

**Conclusions.** Our observations support the scenario in which the dominant effect of the dense  $z \approx 1.6$  cluster environment is an accelerated mass-assembly timescale ( $\sim 1$  Gyr or shorter) through merging activity that is responsible for driving core galaxies across the mass-quenching threshold of  $\log(M_*/M_{\odot}) \approx 10.4$ . Beyond this mass limit, star formation is suppressed on timescales of  $\sim 1$  Gyr, while the direct environmental quenching process seems to be subdominant and is acting on significantly longer timescales ( $\sim 2\text{--}3$  Gyr).

**Key words.** X-rays: galaxies: clusters – galaxies: clusters: individual: XDCP J0044.0-2033 – galaxies: elliptical and lenticular, cD – galaxies: evolution – galaxies: formation – galaxies: luminosity function, mass function

## 1. Introduction

The discovery of clusters of galaxies at redshifts  $1.5 < z \lesssim 2$  over the past few years with infrared (e.g., Gobat et al. 2011; Zeimann et al. 2012; Stanford et al. 2012; Muzzin et al. 2013; Papovich et al. 2010; Andreon et al. 2014) or X-ray selection techniques

(e.g., Fassbender et al. 2011c; Santos et al. 2011; Tozzi et al. 2013; Tanaka et al. 2013) has enabled a new view on galaxy evolution studies in what was formerly known as the “redshift desert” with large ensembles of co-eval galaxies in the densest environments. However, owing to the detection challenge and the relative rareness of these systems, detailed studies of the galaxy population properties in the densest  $z > 1.5$  environments have so far only been available for a small number of galaxy groups with total masses of  $M_{200} < 10^{14} M_{\odot}$  (e.g., Rudnick et al. 2012; Papovich et al. 2012; Strazzullo et al. 2013; Newman et al. 2014).

While the galaxy populations in the centers of massive galaxy clusters show surprisingly little change in their overall

\* Based on observations under programme ID 084.A-0844, 087.A-0351, and 089.A-0419 collected at the European Organisation for Astronomical Research in the Southern Hemisphere, Chile.

\*\*  $J$ - and  $K_s$ -band FITS files are available at the CDS via anonymous ftp to [cdsarc.u-strasbg.fr](http://cdsarc.u-strasbg.fr) (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/568/A5>

appearance and properties all the way out to  $z \simeq 1.4$  (e.g., Mei et al. 2009; Strazzullo et al. 2010), it has now been firmly established that the  $z \simeq 1.5$  regime is a crucial transition epoch beyond which the predecessors of the homogeneous “red and dead” galaxy population of lower- $z$  systems undergo drastic changes as the star formation activity is switched on in many systems all the way to the densest core regions (e.g., Tran et al. 2010; Hayashi et al. 2010; Hilton et al. 2010; Fassbender et al. 2011c; Bayliss et al. 2013; Romeo et al. 2008).

However, observations of high star formation rates in cluster cores and active mass-assembly activity (e.g., Lotz et al. 2013) in massive group galaxies and the associated gradual disappearance of a well-defined red sequence on one hand are currently confronted with recent reports on seemingly evolved galaxy populations in some systems at  $z \simeq 1.6$  (e.g., Tanaka et al. 2013) and beyond (Andreon et al. 2014). Current observations of the general evolutionary state of galaxies in  $z \gtrsim 1.5$  group and cluster environments therefore do not provide a homogeneous picture yet owing to the still limited statistics of investigated systems and, in particular, the lack of mass leverage into the more massive cluster regime.

Clearly, the influence and effects of the total cluster mass  $M_{200}$  on the overall galaxy evolution process at  $z \gtrsim 1.5$  is one missing key aspect to be investigated further. Moreover, at lookback times of around 9.5 Gyr and beyond, the cluster formation and galaxy evolution timescales become similar to the cluster ages since gravitational collapse. Therefore the individual mass accretion histories of the cluster can be expected to play a significant role in the evolutionary state of the overall system and its galaxy populations. Improved sample statistics and a wide mass-baseline at  $z \gtrsim 1.5$  are therefore key to any detailed understanding of the early formation and evolution process of group and cluster galaxies.

In this study we aim to investigate the galaxy population properties of the massive cluster XDCP J0044.0-2033 at  $z = 1.58$  (Santos et al. 2011) using new deep near-infrared (NIR) observations in the  $J$ - and  $Ks$  bands along other complementary data sets. The cluster was originally discovered based on its extended X-ray emission (centroid coordinates: RA = 00:44:05.2, Dec = -20:33:59.7) within the framework of the *XMM-Newton* Distant Cluster Project (XDCP, Fassbender et al. 2011a), which implies that the selection is unbiased with respect to the underlying galaxy population of the system. Based on the discovery X-ray data set, the cluster has a measured bolometric luminosity of  $L_{X,500}^{\text{bol}} \simeq 6 \times 10^{44} \text{ erg s}^{-1}$  and a derived luminosity-based total system mass estimate (Reichert et al. 2011) of  $M_{200} \simeq 3 \times 10^{14} M_{\odot}$ , making it the most massive X-ray selected cluster at  $z > 1.5$  currently known. This paper is the first in a series of upcoming studies with the aim to provide a very detailed multi-wavelength characterization of the system. These works will include, for example, updated cluster mass measurements based on new observations with *Chandra* (Tozzi et al., 2014), the total cluster star formation activity derived from *Herschel* data (Santos et al., in prep.), and studies based on *Spitzer* observations and additional spectroscopic and optical imaging coverage.

The structure of this paper is as follows: in Sect. 2 we introduce the imaging and spectroscopic data sets used for this study and discuss the data reduction; in Sect. 3 we present the results of various aspects of the galaxy population properties of the system. Section 4 provides a global view of the cluster, Sect. 5 discusses the findings and compares them with galaxy evolution models, and we conclude in Sect. 6. Throughout this paper we assume a standard  $\Lambda$ CDM cosmology with  $\Omega_m = 0.3$ ,  $\Omega_{\Lambda} = 0.7$ , and  $h = H_0/(100 \text{ km s}^{-1} \text{ Mpc}^{-1}) = 0.7$ . At a redshift of  $z = 1.58$ ,

this implies a lookback time of 9.45 Gyr, an age of the Universe of 4.0 Gyr, and an angular scale of  $8.47 \text{ kpc}''$  very close to the absolute maximum of the angular diameter distance  $D_A$ . We use the notation with subscript  $X_{500}$  ( $X_{200}$ ) to refer to physical quantities  $X$  measured inside a radius, for which the mean enclosed total density of the cluster is 500 (200) times the critical energy density of the Universe  $\rho_{\text{cr}}(z)$  at the cluster redshift. All reported magnitudes are given in the Vega system.

## 2. Data sets and data reduction

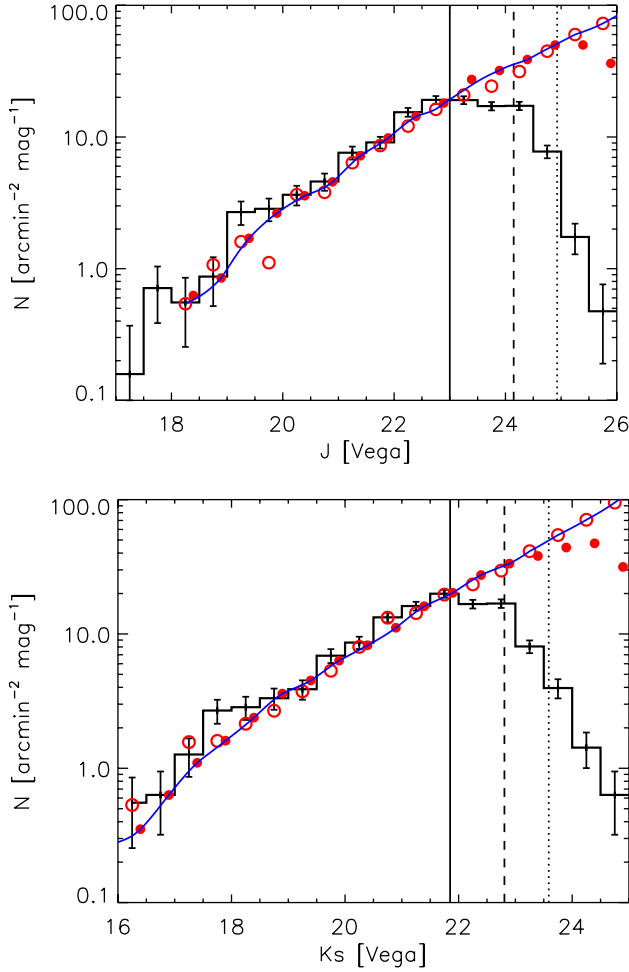
The data sets for this work include deep NIR imaging observations in the  $J$ - and  $Ks$  band with VLT/HAWK-I presented in Sect. 2.1, complemented by new spectroscopic data (Sect. 2.2), archival Subaru imaging in the  $V$ - and  $i$  band (Sect. 2.3), and mid-infrared observations with *Spitzer* (Sect. 2.4).

### 2.1. Near-infrared observations with VLT/HAWK-I

The cluster field of XDCP J0044.0-2033 was observed with VLT/HAWK-I (Kissler-Patig et al. 2008) on 27 June 2011 and 08 August 2011 with the  $J$ - and  $Ks$ -band filters under excellent observing conditions with  $0.5''$  seeing (program ID 087.A-0351). In total, the field was observed for a net exposure time of 80 min in  $J$  (40 positions each with  $12 \times 10 \text{ s}$ ) and 36 min in  $Ks$  (36 positions each with  $6 \times 10 \text{ s}$ ) with slight telescope dithering offsets between observing positions. The focal plane of HAWK-I consists of four  $2048 \times 2048$  pixel detectors with  $15''$  gaps between them, providing a field of view (FoV) for each quadrant of  $3.6'$  on the side with a pixel scale of  $0.106''/\text{pixel}$ . The cluster center was placed well inside one quadrant (Q3) to provide a homogeneous deep exposure time in the cluster region without interference of the detector gaps. For this study we used the other three quadrants as control fields to accurately determine the background source density at the full depth of the observations.

The data reduction was performed in standard NIR manner, closely following the processing steps described in Lidman et al. (2008). In brief, these steps include dark-frame subtraction, flat-fielding, background subtraction with object masking, astrometric calibration, and the final seeing-weighted image combination using *SWarp* (Bertin et al. 2002). The seeing of the resulting deep combined images was measured to be  $0.51''$  (FWHM) in  $Ks$  and  $0.53''$  in the  $J$  band. The photometric zero points for both bands were determined to an accuracy of about 0.02 mag (0.05 mag) in  $J$  ( $Ks$ ) using 2MASS point sources (Skrutskie et al. 2006) in the FoV as well as dedicated standard star observations.

The photometry was performed with *SExtractor* (Bertin & Arnouts 1996) in dual-image mode. Since the depth of the  $J$ - and  $Ks$ -band image relative to the expected magnitudes of the red cluster galaxy population was designed to be comparable, we co-added the final image stacks, which resulted in a noise-weighted combined  $J + Ks$  detection image with the maximum achievable depth. This combined frame was then used as object-detection image for the first *SExtractor* pass, whereas the final photometry was measured in the original individual bands in  $J$  and  $Ks$ . To accurately determine object colors the  $Ks$  image quality was matched to the slightly different seeing conditions in the  $J$  image. The magnitudes of the resulting photometric catalogs were corrected for galactic extinction. We used MAG\_AUTO for the total object magnitudes and  $1''$  aperture magnitudes ( $\sim 2 \times$  seeing) to determine the  $J - Ks$  object colors. Finally, we applied the small color term correction to transform



**Fig. 1.** Differential galaxy number counts in the  $J$  (top panel) and  $K_s$ -band (bottom panel). The black histograms show the measured differential galaxy counts in the background regions with Poisson errors. Open circles show literature number counts taken from Maihara et al. (2001) while filled circles are counts from larger fields from Windhorst et al. (2011) in  $J$  and Galametz et al. (2013) in  $K_s$ . The solid lines show the combined reference counts used to determine the completeness levels as a function of magnitude. The 100%/50%/10% completeness levels are indicated by solid/dashed/dotted vertical lines. The slight excess in some bins can be attributed to the large-scale structure environment of XDCP J0044.0-2033 and the nearby cluster A2813 at  $z = 0.292$ .

the HAWK-I into the 2MASS filter system following Lidman et al. (2008).

Figure 1 shows the measured differential galaxy number counts in three dedicated background regions, that is, in the three quadrants that do not contain the cluster center. These background regions with a combined sky area of 25.3 square arcmin were selected as areas with the full homogeneous exposure time and without contaminating bright stars. The measured galaxy counts (black histograms) are compared with published literature counts (Maihara et al. 2001; Windhorst et al. 2011; Galametz et al. 2013) in the same filters to evaluate the completeness levels of the observations as a function of limiting magnitudes. These limiting magnitudes for completeness levels of 100%/50%/10% (vertical lines) were determined to be 23.0/24.2/24.9 mag (21.9/22.8/23.6 mag) in the  $J$  ( $K_s$ ) band.

In some of the bins, in particular for brighter magnitudes, a slight excess of observed counts on the  $1-2\sigma$  level with respect the reference counts is observed. This is to be expected

given the proximity of the observed HAWK-I field to the massive low- $z$  cluster A2813 at  $z = 0.292$ , which was the target of the original *XMM-Newton* archival field. The mass of A2813 of  $M_{200} \approx 7 \times 10^{14} M_\odot$  translates into a cluster radius of  $R_{200} \approx 1.7$  Mpc  $\approx 6.6'$  (Zhang et al. 2008). The distant cluster XDCP J0044.0-2033 is located approximately 10.5' to the north-east of A2813, implying that the HAWK-I background fields are at cluster-centric distances of  $1.5-2 R_{200}$  away from the A2813 center, that is, still in the large-scale-structure (LSS) environment for which a cluster-associated excess of foreground galaxies can be expected. Moreover, the background-field locations also are at similar cluster-centric distances in units of  $R_{200}$  from XDCP J0044.0-2033 itself, implying that the LSS of the system under investigation might contribute in a non-negligible way to the observed background counts. The appropriate background reference counts, that is, measured in background fields versus literature counts, to be used thus depend on the specific applications discussed in Sect. 3.

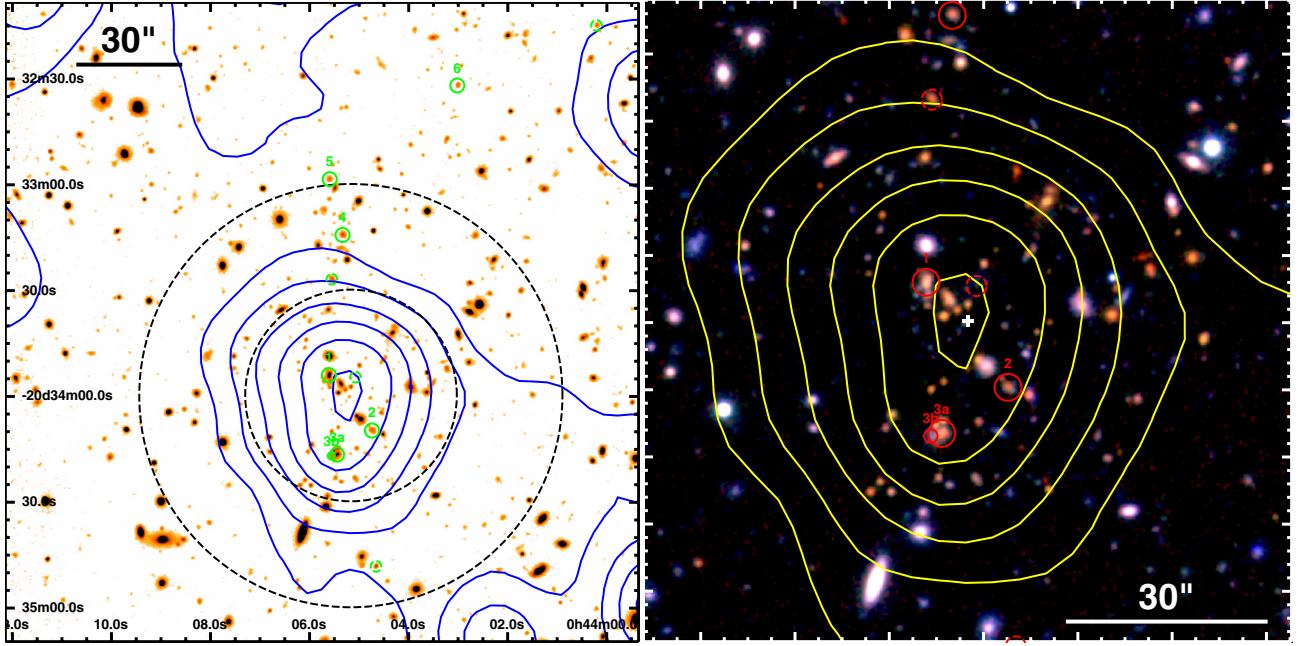
## 2.2. Optical spectroscopy with VLT/FORS 2

The original discovery data set of XDCP J0044.0-2033 presented in Santos et al. (2011) included only three secure spectroscopic member galaxies of the cluster. To increase the number of spectroscopically confirmed members, we initiated a new optical spectroscopic campaign with VLT/FORS 2 for a new deep MXU (mask exchange unit) observation using a customized mask of targeted color-selected candidate member galaxies (program ID: 089.A-0419). However, during the recent observing season only one-third of the scheduled observations was completed, resulting in a net on-target exposure time of 1.45 h in good subarcsec observing conditions executed on 1 and 2 January 2013.

The data were reduced in standard manner using the new publicly available FORS 2 pipeline F-VIPGI described in detail in Nastasi et al. (2013). The limited exposure time allowed only the secure redshift determination for member galaxies with significant [O II]-emission line. This way, we identified four new secure emission-line cluster members listed in Table 1 with their identification numbers (2, 3a, 5, and 6) together with the redshifts, their cluster-centric distance, total  $K_s$  magnitude, and  $J - K_s$  and  $i - K_s$  colors. Together with the previously known members (IDs 1, 3b, and 4) this increases the number of known secure spectroscopic members to seven, leaving the previously determined median system redshift unchanged at  $z_{\text{med}} = 1.579$ . Additionally, four new tentative cluster members were identified, whose redshifts await a final confirmation once the full observational data set is available. The spatial distribution of all secure and tentative spectroscopic cluster member galaxies are indicated in Fig. 2. More details on the spectroscopic properties of confirmed cluster member galaxies will be presented in a forthcoming publication for a combined analysis of all new spectroscopic data sets (Nastasi et al., in prep.).

The double object with IDs 3a and 3b is of particular interest. From its NIR morphology it is identified as a single object (3a) by *SExtractor*, see also the upper right panel of Fig. 10 for a close-up view. However, the subcomponent 3b has a significantly bluer color than the central part of the main component 3a, which establishes their nature as two distinct objects, as can be seen in Fig. 17. This is now confirmed by the new spectroscopic results, which yield two distinct spectra and redshifts with  $z = 1.5699$  for the main component 3a and  $z = 1.5716$  for the bluer subcomponent 3b. With a projected separation of about  $1.6'' \approx 13$  kpc and a restframe velocity offset of  $\Delta v \approx 200$  km s $^{-1}$  we can establish this double system as an ongoing merging event





**Fig. 2.** Environment of the galaxy cluster XDCP J0044.0-2033 at  $z = 1.58$ . Secure spectroscopic cluster members are marked by solid small circles with their object ID number on top, the positions of tentative system members are indicated by small dashed circles. The logarithmically spaced contours show the spatial distribution of the X-ray emission detected with *XMM-Newton*. *Left panel:*  $3' \times 3'$  overview image of the combined  $J + Ks$  detection band including all currently known secure and tentative cluster member galaxies. The large dashed circles indicate the  $30''$  and  $60''$  ( $\approx R_{500}$ ) radii about the X-ray centroid position. *Right panel:* color composite image centered on the X-ray centroid (white cross) made up of the  $Ks$  (red channel),  $J$  (green), and  $i$ -band (blue) with dimensions  $1.6' \times 1.6'$ .

of the blue satellite galaxy (3b) with the main red galaxy (3a), which is the second-ranked member galaxy in NIR luminosity among the secure members.

### 2.3. Subaru $V$ - and $i$ -band data

The field of the low- $z$  cluster A2813 has archival coverage of good quality wide-field imaging data in the  $V$ - and  $i$  bands with Subaru/Suprime-Cam originally intended for a weak-lensing study of A2813. These data were taken on 6 November 2010 in subarcsec observing conditions for a total exposure time of 28 min ( $7 \times 240$  s) in  $V$  and 36 min ( $9 \times 240$  s) in  $i$ . These archival data complement our dedicated follow-up programs of XDCP J0044.0-2033 in the near- and mid-infrared toward the optical bands with similar depth and image quality.

The Subaru data were reduced in standard manner using the SDFRED2 software (Ouchi et al. 2004). The resulting image quality of the final stacked images was measured to be  $0.77''$  (FWHM) in  $V$  and  $0.72''$  in the  $i$  band. Photometric zero points were obtained from observations of the standard-star field SA95 (Landolt 1992) during the same night. The optical Subaru images were co-aligned with the HAWK-I data and re-gridded to the same pixel scale. Photometric catalogs were obtained in the same manner as described in Sect. 2.1 with appropriate seeing adjustments for color measurements in  $i - Ks$ . The limiting magnitudes for the  $i$ -band image were determined to be  $27.0/28.0/28.5$  mag (Vega) for completeness levels of 100%/50%/10%.

### 2.4. Mid-infrared observations with *Spitzer*/IRAC

*Spitzer*/IRAC observations in the  $3.6 \mu\text{m}$  and  $4.5 \mu\text{m}$  channels were obtained for XDCP J0044.0-2033 on 24 August 2011 as part of program 80136. The total integration time in each channel

was 270 s tailored to reach a depth for the detection of individual galaxies of up to  $m^* + 2$  at  $z \sim 1.6$ .

We retrieved processed (post-BCD) images for both channels from the *Spitzer* Heritage Archive. Tests were performed with the APEX/MOPEX pipeline to assess the quality of the reduction, which was found to be sufficient for our science purposes without additional processing steps. The photometry of individual objects was performed with *SExtractor* in both channels, which for this work was used only to mask bright foreground galaxies.

For the current work we only aim at obtaining a measurement of the integrated  $4.5 \mu\text{m}$  MIR flux within a given projected distance from the cluster center independent of the detection and proper deblending of individual galaxies, for which we applied the following approach: we masked out foreground sources brighter than the most luminous cluster galaxies with  $m(\text{Vega}) \leq 15$  mag (see left panel of Fig. 6). We then measured the average integrated background surface brightness for the  $4.5 \mu\text{m}$  channel in a dozen independent background regions outside the cluster volume, most of which are located in the parallel field observations obtained simultaneously during the channel 1 on-target integration. This determination of the integrated  $4.5 \mu\text{m}$  background surface brightness yielded a field value of  $0.1378 \pm 0.0028$  MJy/sr with a standard deviation representative for square arcmin scales. This background value was subsequently subtracted from the integrated  $4.5 \mu\text{m}$  flux measured inside the analysis region of XDCP J0044.0-2033.

## 3. Galaxy population properties of XDCP J0044.0-2033

Figure 2 provides an overview of the cluster and its environment together with the spatial distribution of the detected *XMM-Newton* X-ray emission and the location of the

**Table 1.** Spectroscopically confirmed, secure member galaxies of the cluster XDCP J0044.0-2033.

ID	RA J2000	Dec J2000	$z_{\text{spec}}$	$d_{\text{cen}}$ arcsec	$Ks$ Vega mag	$J - Ks$ Vega mag	$i - Ks$ Vega mag
1	11.02336	-20.56509	1.5797	8.2	17.96	1.98	4.21
2	11.01973	-20.56936	1.5936	11.7	19.52	2.08	4.23
3a	11.02269	-20.57125	1.5699	17.3	18.06	2.33	4.86
3b	11.02316	-20.57135	1.5716	18.0	21.04	1.38	2.98
4	11.02220	-20.55395	1.5787	45.5	19.25	2.10	4.27
5	11.02333	-20.54957	1.5778	61.6	19.58	1.92	4.20
6	11.01255	-20.54218	1.5986	92.9	20.09	1.82	3.71

**Notes.** The table lists the object IDs, the galaxy coordinates, the cluster-centric distances  $d_{\text{cen}}$  relative to the *XMM-Newton* determined centroid of the extended X-ray emission, the total  $Ks$ -band magnitude, the  $J - Ks$  color, and the  $i - Ks$  color within  $1''$  apertures. Typical measurement uncertainties for the individual galaxy redshifts are  $\sigma_z \sim 0.0005$ . The  $Ks$  magnitude of ID 3b is given for a  $1''$  aperture since there is no clear separation from the main component 3a (see upper right panel of Fig. 10). All galaxies but ID 1 exhibit a detected [O II] emission line.

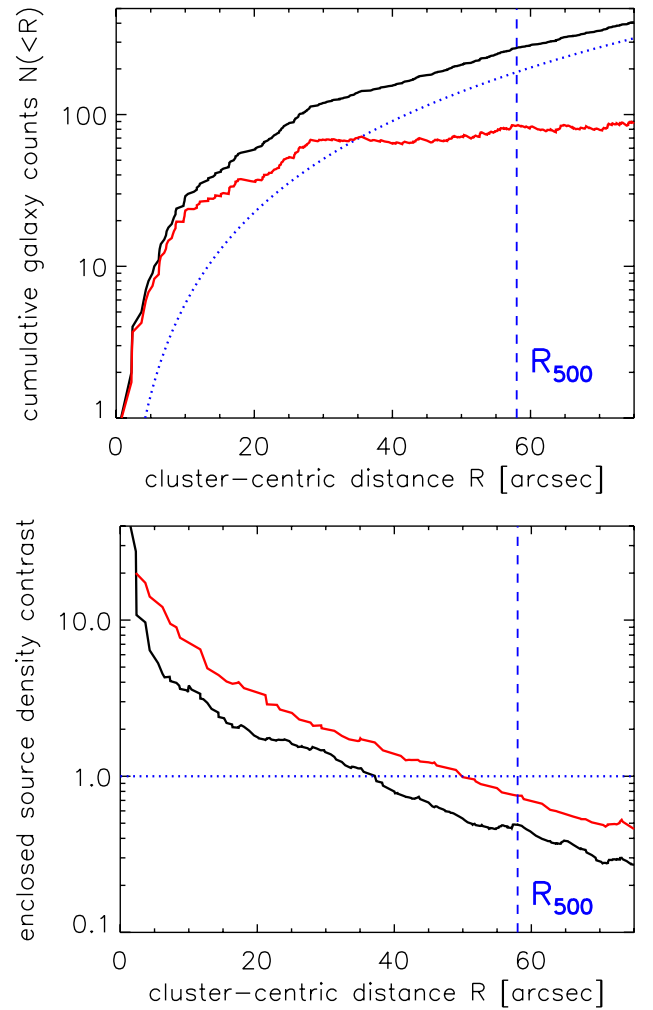
spectroscopic members. The image in the left panel covers an area of  $3' \times 3'$ , while the  $i + J + Ks$  band color composite in the right panel shows the central  $1.6' \times 1.6'$  of the galaxy cluster. From the original total mass estimate of  $M_{200} \simeq 3 \times 10^{14} M_{\odot}$  we can approximate the characteristic cluster radii to be  $R_{200} \simeq 760 \text{ kpc} \simeq 90''$  and  $R_{500} \simeq 490 \text{ kpc} \simeq 58''$ .

In the following subsections, we investigate different aspects of the cluster galaxy population properties starting with purely statistical measures such as radial surface density profiles (Sect. 3.1), the stellar mass profile (Sect. 3.2), and NIR luminosity functions (Sect. 3.3), followed by a closer look at the morphologies of individual galaxies (Sect. 3.4), the color–magnitude relation (Sect. 3.5), the faint end of the red sequence (Sect. 3.6), the spatial distribution of color–color selected galaxy types (Sects. 3.7 and 3.8), and the formation of the brightest cluster galaxy in the core region (Sect. 3.9).

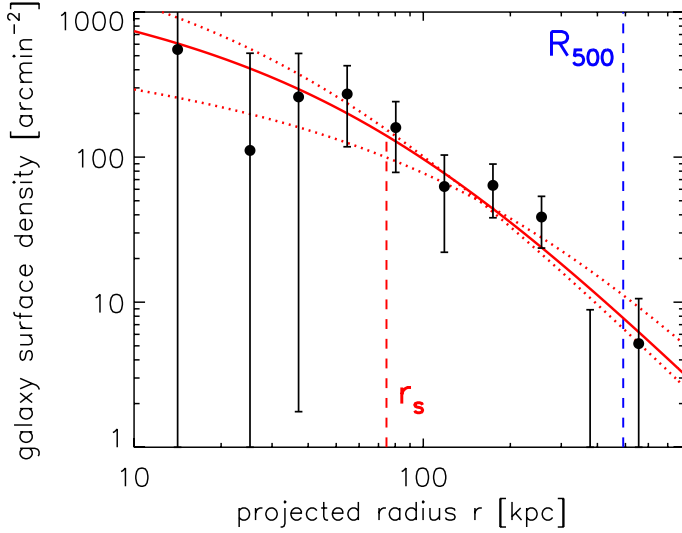
### 3.1. Radial galaxy density profiles

As a first application of the deep NIR HAWK-I data of Sect. 2.1, we measured the overdensity contrast and radial profile of the galaxy population of XDCP J0044.0-2033. To this end, we determined the background galaxy density of all nonstellar sources in the three control quadrants as  $68.1 \pm 2.5 \text{ arcmin}^{-2}$ , where the background uncertainty is given by Poisson error along with the expected influence of the LSS contribution of the cluster environments. The top panel of Fig. 3 shows the detected cumulative galaxy counts  $N(<R)$  within the cluster-centric distance  $R$  for all galaxies, the expected background, and the net excess of cluster galaxies. In total, we detect a cluster associated net excess of about 90 galaxies, approximately 75% of which are located within 30 arcsec from the cluster center. The enclosed high- $z$  cluster galaxy population equals the background galaxy contribution at a cluster-centric distance of 35 arcsec, outside of which the background is dominant.

The enclosed density contrast of cluster galaxies with respect to the background is shown in the bottom panel of Fig. 3 for the bulk of the galaxy population and bright galaxies alone. At  $z \sim 1.6$ , even a rich cluster such as XDCP J0044.0-2033 stands out from the background only within about half an arcmin in deep NIR imaging data. Source-density contrasts on the order of 10 or more are only reached within a few arcsec from the X-ray center. When the enclosed density profile is restricted to brighter magnitudes  $17 \leq Ks \leq 20$  (red line), for instance, with shallower NIR data, the density contrast is shifted upward owing to a significantly reduced background density ( $\simeq 10.5 \text{ arcmin}^{-2}$ )



**Fig. 3.** Enclosed galaxy count properties within the projected cluster-centric distance  $R$  from the X-ray centroid based on all HAWK-I detected sources. The dashed vertical line indicates the  $R_{500}$  radius of the cluster. *Top panel:* cumulative radial galaxy count profile  $N(<R)$  of enclosed sources within  $R$  for all galaxies (black solid line) and background-subtracted net cluster member counts (red line). The average background counts are shown by the blue dotted line. A total of about 90 excess galaxies associated with the cluster are detected. *Bottom panel:* enclosed galaxy density contrast  $([N(<R) - B(<R)]/B(<R))$  within radius  $R$  for all galaxies with  $17 \leq Ks \leq 22$  (black line) and brighter galaxies with  $17 \leq Ks \leq 20$  (top red line). At the intersection point with the horizontal dashed dotted line the enclosed cluster galaxy density equals the background density.

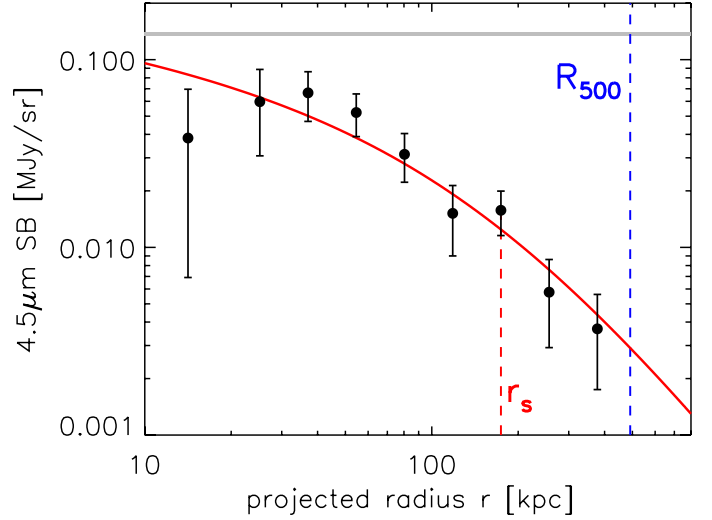


**Fig. 4.** Projected cluster galaxy surface-density profile (black data points) and best-fitting projected NFW profile (red solid line) with a concentration parameter of  $c_{200} = 10.2 \pm 6.6$  and scale radius  $r_s = R_{200}/c_{200} \approx 75$  kpc (dashed red vertical line). Dotted red lines depict the profiles with  $c_{200} \pm 1\sigma$ .

and the preferred location of luminous cluster members in the central region.

With the given number of detected cluster-associated galaxies for XDCP J0044.0-2033 one can attempt to measure the NIR galaxy surface-density profile of the cluster. This is shown in Fig. 4 as a function of projected physical cluster-centric radius with Poisson errors following Gehrels (1986). The red solid line shows the best fit projected NFW profile (e.g., Lokas & Mamon 2001) to the source surface density distribution with a formal best-fit solution for the concentration parameter of  $c_{200} = 10.2 \pm 6.6$ . Despite the significant statistical uncertainty in determining the concentration parameter, as can be expected for a single galaxy cluster at  $z \sim 1.6$ , the compact distribution of cluster galaxies can be visually confirmed in the color image of Fig. 2 (right panel), where very few red galaxies are visible beyond  $R \gtrsim 30''$ , in particular toward the south (bottom) and east (left).

Based on the compact best-fit galaxy surface-density profile with  $c_{200} \approx 10$ , constrained mostly within  $R_{500}$ , with its relatively high concentration parameter compared to local clusters (e.g., Budzynski et al. 2012), estimates for the expected number of detectable galaxies in the cluster outskirts can be obtained. This suggests that only about 20% of all galaxies within the fiducial cluster radius  $R_{200}$  are located in the cluster outskirts beyond the  $R_{500}$  radius, that is, at cluster-centric distances of 58–90''. However, in these outer regions the Poisson noise of the background counts dominates the modest expected increase of the cumulative cluster-galaxy count function (top panel of Fig. 3). With the current data, this function can be reliably traced out to a cluster-centric radius of  $\approx 75''$ , beyond which the chip gaps of the HAWK-I detector array influence the source detection. However, the almost flat cumulative profile shape in the outer cluster region with only a modest increase in excess cluster-counts is fully consistent with the extrapolated expectations from the best-fit NFW surface density function within the uncertainties imposed by the increasing Poisson noise.



**Fig. 5.**  $4.5 \mu\text{m}$  surface-brightness profile (black data points) based on the *Spitzer* data shown in Fig. 6. The best-fitting projected NFW profile (red solid line) with a concentration parameter of  $c_{200} = 4.4 \pm 1.6$  is overlotted in red. The median background level is indicated by the gray horizontal band on top, vertical lines are as in Fig. 4.

### 3.2. Mid-infrared view of the cluster and stellar mass profile

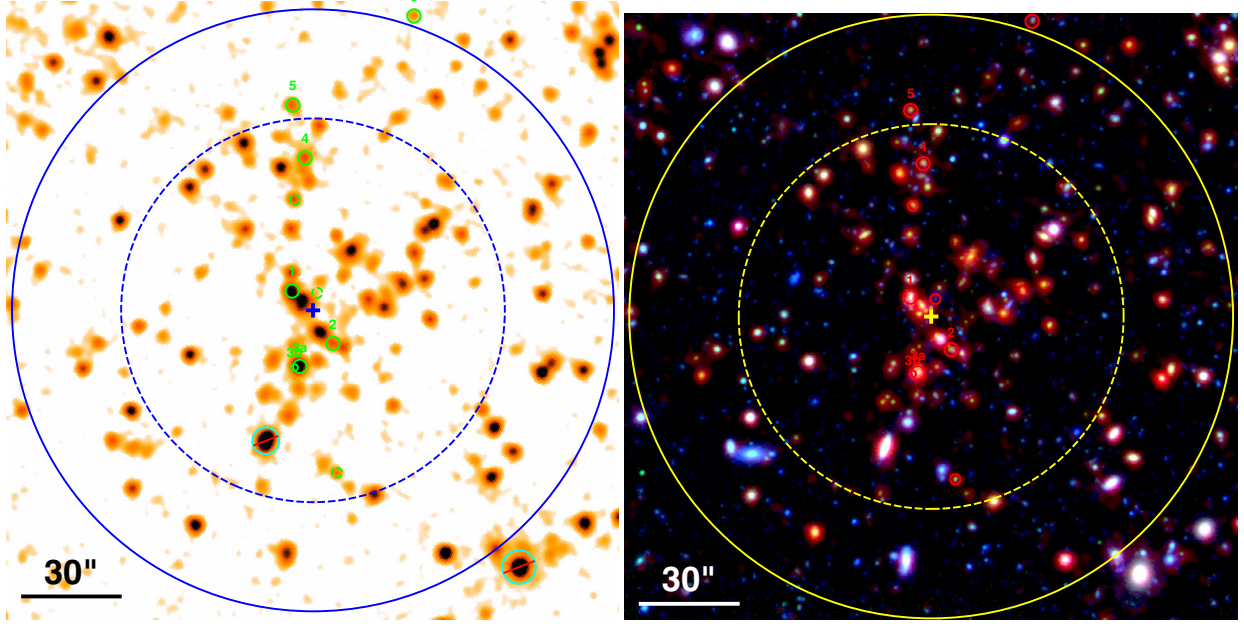
After constraining the radial number-density profile of the cluster-associated galaxy population of XDCP J0044.0-2033, we analogously examined the stellar mass distribution of the cluster. To this end, we used the mid-infrared (MIR) *Spitzer*/IRAC observations introduced in Sect. 2.4. At  $z \sim 1.6$ , the observed  $4.5 \mu\text{m}$ -band traces the restframe NIR peak of stellar light emitted by old stellar populations (see Fig. 11) and is therefore optimally suited to derive integrated quantities of the stellar light and mass content of high- $z$  clusters (e.g., Brodwin et al. 2006).

The *Spitzer*/IRAC view of the projected cluster volume of XDCP J0044.0-2033 is shown in Fig. 6 for the  $4.5 \mu\text{m}$  surface brightness distribution (left panel) and as a color-image representation in combination with the NIR and optical data (right panel). Using the determined integrated background value of Sect. 2.4, we measured a total net  $4.5 \mu\text{m}$  cluster flux of  $1.78 \pm 0.40$  mJy ( $2.58 \pm 0.62$  mJy) inside the projected  $R_{500}$  ( $R_{200}$ ) radius shown in Fig. 6, corresponding to a total integrated apparent (Vega) magnitude of  $m_{500}^{4.5 \mu\text{m}} = 12.45 \pm 0.22$  mag and  $m_{200}^{4.5 \mu\text{m}} = 12.04 \pm 0.24$  mag, respectively.

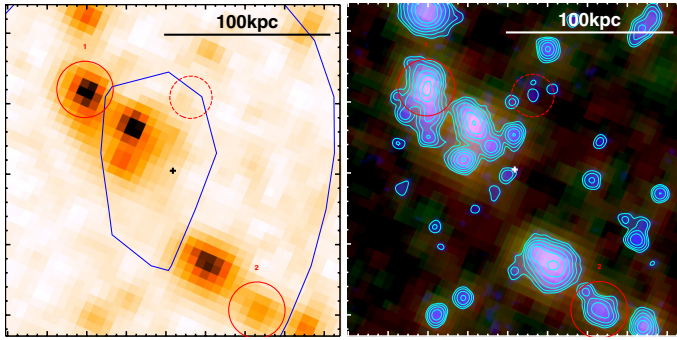
Figure 5 shows the observed  $4.5 \mu\text{m}$  surface brightness profile inside  $R_{500}$ . The best-fitting projected NFW profile to the observed surface density yields a result for the concentration parameter of  $c_{200} = 4.4 \pm 1.6$ , lower than the compact profile parameter of the galaxy surface density of Fig. 4. In contrast to the galaxy number density, the  $4.5 \mu\text{m}$  surface brightness, used as proxy for stellar mass, is not centrally peaked, but drops within  $\approx 40$  kpc of the X-ray centroid position.

The central 100 kpc region of the cluster core is shown as a zoom in Fig. 7 and is discussed in more detail in Sect. 3.9. The brightest central *Spitzer* source is located northeast of the X-ray center (left panel) and is a blend of at least four red galaxies closest to the center (see right panel). The most central galaxies at the current accuracy of the centroid position are faint blue galaxies with little flux in the *Spitzer* band, but counting toward the galaxy surface density with a resulting higher NFW concentration parameter. However, outside the very core region the NFW profile fits the observed  $4.5 \mu\text{m}$  surface brightness profile





**Fig. 6.** *Left panel:* *Spitzer*/IRAC 4.5  $\mu\text{m}$  view of the cluster volume ( $3' \times 3'$ ). The cluster center (central cross),  $R_{500}$  (dashed blue circle), and  $R_{200}$  (solid blue circle) are indicated, small circles mark spectroscopic members as in Fig. 2. The crossed-out cyan objects were removed from the analysis. *Right panel:* color composite with the same FoV by adding the  $V$ -band (blue) and the combined  $J + K_s$  image (green) to the *Spitzer* 4.5  $\mu\text{m}$  data (red).

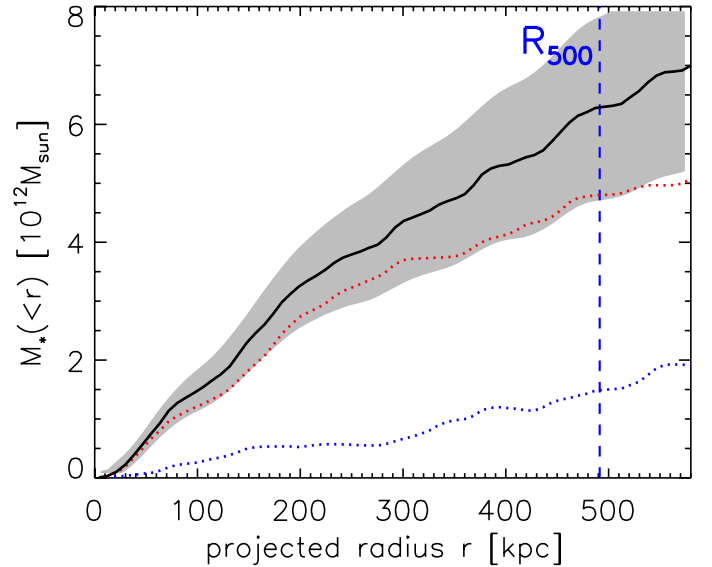


**Fig. 7.** *Spitzer*/IRAC view of the central 100 kpc of the cluster core of XDCP J0044.0-2033. *Left panel:* 4.5  $\mu\text{m}$  image with X-ray contours overlaid in blue and symbols as in Fig. 6. *Right panel:* same field-of-view as color image composed of 4.5  $\mu\text{m}$  (red), 3.6  $\mu\text{m}$  (green), and the 0.5''-resolution NIR  $JK_s$  detection image (blue) with cyan NIR isophotal contours derived from the latter band to visualize the source confusion at the *Spitzer* resolution.

quite well, which is also reflected in the reduced  $\chi^2$  value of the fit of 0.8.

To convert the observed 4.5  $\mu\text{m}$  cluster flux to stellar mass at  $z = 1.58$ , we evaluated the stellar-mass-to-light-ratio  $M_*/L_{4.5 \mu\text{m}}$  based on a grid of stellar population models with different star formation histories and metallicities. To account for changes of  $M_*/L_{4.5 \mu\text{m}}$  across the cluster volume owing to the varying underlying galaxy populations, we adopted a two-model approach that accounts for old red galaxies with a solar metallicity  $z_f = 3$  model on one hand and bluer galaxies with a younger ( $z_f = 2$ ) stellar population on the other. To this end, we masked the areas of detected galaxies with red colors of  $J - K_s \geq 1.84$  (see Sects. 3.5 and 3.6) and converted the background-subtracted cluster flux with the  $z_f = 3$  model, while using the  $z_f = 2$  model for the unmasked regions.

Figure 8 shows the resulting cumulative total stellar mass profile enclosed within the projected cluster-centric radius  $r$  and



**Fig. 8.** Cumulative total stellar mass profile (black solid line) for XDCP J0044.0-2033 as a function of projected cluster-centric radius with  $1\sigma$  uncertainties indicated by the gray-shaded area. The relative contributions to the total mass profile of the red-masked objects (red) and the remaining area (blue) are shown by the dotted lines.

the two contributing components for the old galaxy model and the galaxies with an assumed younger stellar formation epoch. The gray-shaded area shows the  $1\sigma$  uncertainties of the total enclosed stellar mass, which are dominated by net cluster flux errors and the remaining model uncertainties of the stellar-mass-to-light ratios.

We measure a 4.5  $\mu\text{m}$ -based total integrated stellar mass of XDCP J0044.0-2033 within the projected  $R_{500}$  and  $R_{200}$  radii of  $M_{*500} \simeq (6.3 \pm 1.6) \times 10^{12} M_\odot$  and  $M_{*200} \simeq (9.0 \pm 2.4) \times 10^{12} M_\odot$ , respectively. The central 100 kpc cluster core region alone contains a stellar mass of  $M_{*100 \text{ kpc}} \simeq (1.5 \pm 0.3) \times 10^{12} M_\odot$ . Based on the original estimate of the total mass



of XDGP J0044.0-2033 with a  $\sim 35\%$  uncertainty, we can obtain a first approximate determination of the stellar mass fraction<sup>1</sup> of a  $z \sim 1.6$  cluster, yielding  $f_{*,500} = M_{*,500}/M_{500} = (3.3 \pm 1.4)\%$ . This value is fully consistent with the best-fit model of stellar mass fractions of low-redshift systems by Giodini et al. (2009), which predicts a stellar mass fraction of  $\approx 3\%$  with 50% intrinsic scatter for the mass of this cluster.

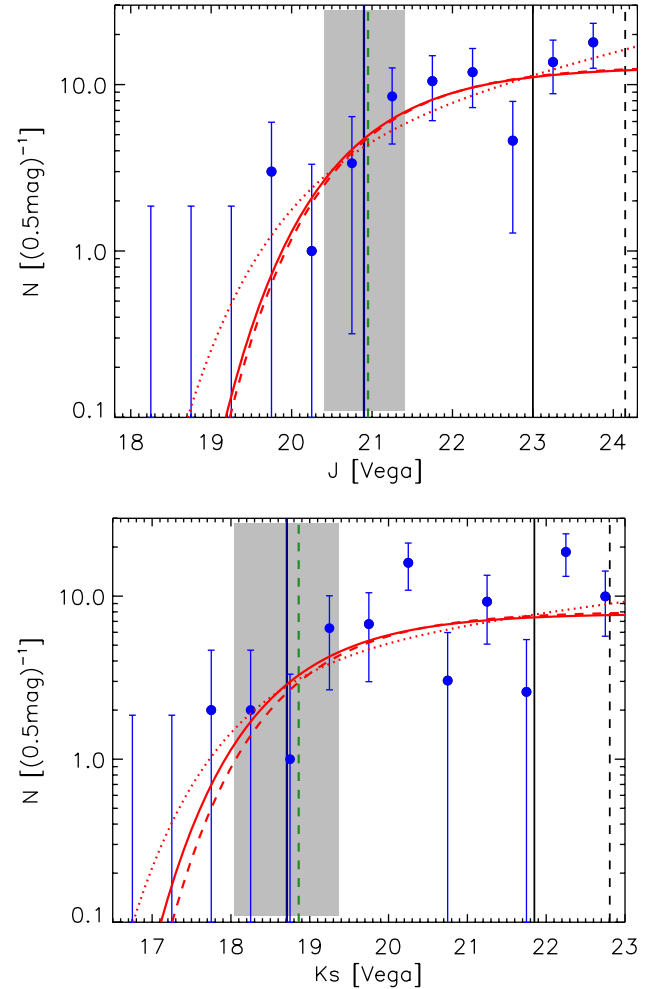
### 3.3. *J*- and *Ks*-band galaxy luminosity functions

Next, we derived the near-infrared galaxy luminosity (LF) functions of XDGP J0044.0-2033 in the *J* and *Ks* bands for a first global characterization of the cluster galaxy population properties. To this end, the signal-to-noise ratio (S/N) of the background-subtracted excess cluster members with respect to the Poisson noise of all galaxies within an encircled region need to be as high as possible. This S/N-optimized cluster-centric extraction radius is determined to be  $30''$  ( $\approx 250$  kpc at  $z = 1.58$ ) with an  $S/N \approx 6.2$ , enclosing approximately 68 cluster galaxies and 52 background sources (see Fig. 3).

To determine the NIR cluster galaxy luminosity functions of XDGP J0044.0-2033 we followed the approach adopted by Strazzullo et al. (2006) for other individual high- $z$  systems at  $z \approx 1.2$ . We selected all galaxies in the S/N-optimized extraction region within  $R \leq 30''$  using a binning of 0.5 mag width. We considered all galaxies with magnitudes up to the 50% completeness limits in *J* and *Ks*, as determined in Sect. 2.1, with the appropriate completeness correction factor applied to the two faintest magnitude bins. The expected background contribution to each bin in the analysis region was then estimated based on the literature model counts (Fig. 1) and subtracted from the measured total counts in a given bin. The uncertainties of the net counts were estimated based on the Poisson errors of both the cluster and background counts following Gehrels (1986) for the low count regime.

At magnitudes brighter than 19 in *Ks* and 20.5 in *J*, the background and cluster counts per bin within the analysis area drop to the order of unity and below, at which point the statistical background-subtraction approach is quite uncertain. To improve the reliability of the LF constraints on the bright end, all eight galaxies in the analysis region with *Ks* < 19 were checked individually and assigned a foreground or cluster membership flag. As can be seen in Fig. 17, three of the galaxies (labels f1-f3) can be readily identified as low- $z$  interlopers and were therefore considered as foreground galaxies. Two galaxies (ID 1 and 3a) are spectroscopically confirmed cluster members, and the three remaining ones are flagged as high-probability cluster member candidates (labels c1-c3) with colors encompassed by the two bright spectroscopic members in a region of the color–magnitude diagram with very low background/foreground contamination (see Sect. 3.5).

Figure 9 shows the resulting observed net cluster counts for XDGP J0044.0-2033 per magnitude bin in the *J*- (top panel) and *Ks* band (bottom panel). To obtain constraints on the characteristic magnitude  $m^*$  at  $z \sim 1.6$  a Schechter function (Schechter 1976) with the additional free parameters for the normalization  $\phi^*$  and the faint end slope  $\alpha$  was fitted to the data. However, a three-parameter LF fit for a single cluster at this redshift has

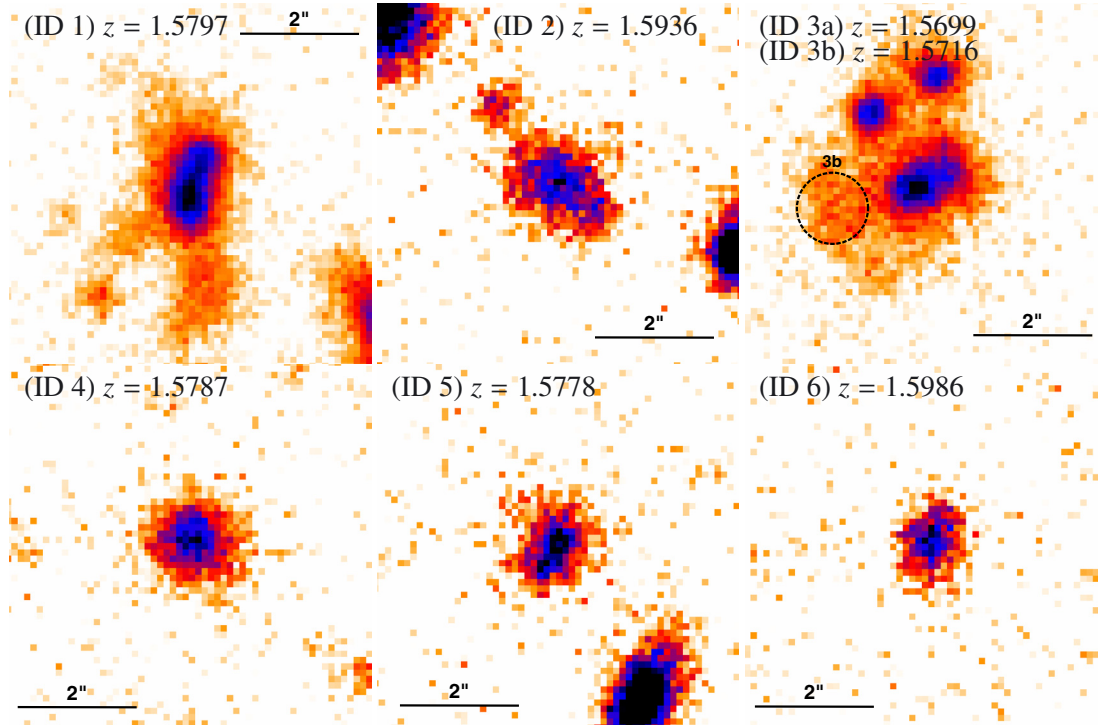


**Fig. 9.** Observed galaxy luminosity functions (blue data points) in the *J*- (top) and *Ks*-band (bottom) for XDGP J0044.0-2033 extracted within a cluster-centric radius of  $30''$  ( $\approx 250$  kpc). The red lines show the best fit Schechter functions for three ( $\phi^*$ ,  $m^*$ ,  $\alpha$ ) free parameters (dotted), two free parameters with a fixed faint end slope of  $\alpha = -1$  (solid), and a single parameter normalization fit (dashed) using a SSP model magnitude for  $m^*$  (green dashed vertical line). The dark blue vertical line with the grey shaded area indicates the best fit solution of the characteristic magnitude  $m^*$  of the two parameter fit with  $1\sigma$  uncertainties. The vertical black (dashed) solid lines on the right indicate the 100% (50%) completeness limits.

too many degrees of freedom for useful constraints on each parameter. As a realistic constrained baseline model we therefore used a two-parameter fit with a fixed faint-end slope of  $\alpha = -1$  as found, for example, by Strazzullo et al. (2006) and Mancone et al. (2012) for their composite cluster LF at  $z \approx 1.2$ . For consistency checks, we also fitted an LF with a single free parameter for the normalization based on the prediction of the characteristic magnitude  $m^*$  from a simple stellar population (SSP) evolution model (Fioc & Rocca-Volmerange 1997).

Although deriving an LF measurement for a single  $z \sim 1.6$  cluster is a challenging task, the observed *J*-band luminosity function of XDGP J0044.0-2033 in the top panel of Fig. 9 is surprisingly well defined and is consistently described by all three underlying Schechter fits with reduced  $\chi^2$  values of  $\sim 0.5$ . The two-parameter fit (solid red line) of the luminosity function yields a characteristic magnitude of  $J^* = 20.90 \pm 0.50$  that agrees very well with the SSP model prediction for a formation redshift of  $z_f = 3$  of  $J_{*SSP} = 20.95$ . The three-parameter

<sup>1</sup> The stellar mass-fraction measurement is to first order independent of the exact value of the total cluster mass, since total mass and stellar mass follow similar radial profiles. However, since the cluster radius changes with total mass, the exact value of  $M_{*,500}$  does increase for higher estimates of the total cluster mass.



**Fig. 10.** Image cutouts of the secure cluster member galaxies based on the combined  $J + K_s$  band detection image without any smoothing applied. Each postage-stamp cutout has dimensions  $6'' \times 6''$  ( $\sim 50$  kpc  $\times$  50 kpc) centered on the member galaxy with object ID and redshift stated on top of each panel. The color scale is linear with increasing flux toward darker colors. The position and redshift of the secondary object clump with ID 3b is indicated in the *upper right panel*.

fit is consistent within the expected large uncertainties ( $J^* = 20.11 \pm 1.60$ ,  $\alpha = -1.30 \pm 0.37$ ).

The LF of the  $K_s$ -band data (bottom panel of Fig. 9) is not quite as well defined as  $J$ , with several outliers on the  $2\sigma$  level from the best-fitting LF models and corresponding reduced  $\chi^2$  values of  $\sim 1.2$ . Nevertheless, the derived characteristic magnitude based on the two-parameter fit of  $K_s^* = 18.71 \pm 0.66$  still agrees well with the SSP model value of  $K_s^*_{\text{SSP}} = 18.86$ . The formal three-parameter fit yields  $K_s^* = 18.17 \pm 1.54$ ,  $\alpha = -1.15 \pm 0.29$ .

Since a single, well-defined  $J - K_s$  color for the cluster galaxies would only result in a shift along the  $x$ -axis between the  $J$  and  $K_s$  data points, one can immediately conclude from the measured magnitude distribution that the cluster galaxies cannot have a single predominant  $J - K_s$  color. This aspect is investigated in more detail in Sect. 3.5 based on the color–magnitude diagram.

### 3.4. Morphologies of secure cluster members

After the purely statistical analyses of the last three sections, we now examine the spectroscopically confirmed secure cluster member galaxies more closely. Although still limited to seven secure members, these galaxies represent a test sample to cross-check and confirm the results of the statistical analyses without the uncertainty introduced by the interloping foreground/background population.

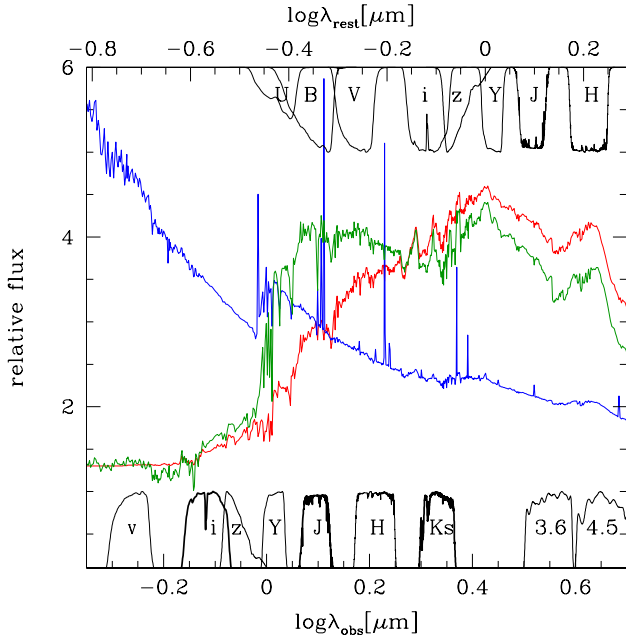
Figure 10 presents  $6'' \times 6''$  postage-stamp cutouts of the secure member galaxies based on the NIR co-added  $J + K_s$  detection image. The corresponding physical projected dimensions of these cutouts are  $50.8$  kpc  $\times$   $50.8$  kpc, with a scale of individual pixels of  $0.9$  kpc/pixel. The seeing-limited resolution scale of

about  $4.5$  kpc ( $0.52''$ ) allows a first-order evaluation of the underlying morphologies. The spectroscopic members in the top row of Fig. 10 are located within the  $30''$  ( $\approx 250$ ) core region of the cluster with the highest densities, whereas members with IDs 4–6 in the bottom row are located at projected radii of  $(0.5 - 1) R_{200}$ , as shown in Figs. 2 and 6 and Table 1.

The most luminous spectroscopic member of the current sample in the NIR is ID 1, which is also a strong candidate for the current brightest cluster galaxy (see Sect. 3.9). This  $m^* - 1$  galaxy is located about  $70$  kpc in projection NE of the X-ray centroid position. Based on its strongly distorted NIR morphology (top-left panel of Fig. 10) in the N-S direction and two further extensions toward the south and SE, it is readily visually classified as system in a state of very dynamic mass-assembly activity with probably several ongoing and recent merger events. Although the color of the main component is bluer than many of the surrounding red galaxies (Figs. 2 and 17), this is the only spectroscopic member without any detected [O II] line emission in the optical spectrum. The two S/SE extensions exhibit even bluer colors than the main component.

The second-most luminous spectroscopic member is the red galaxy with ID 3a with its spectroscopically confirmed smaller blue companion system with ID 3b (see Sect. 2.2). Based on the NIR morphology, this double system is not separable (top right panel of Fig. 10), which is clearly the case when adding the optical bands (see Fig. 17). Overall, the morphology of the main component seems less distorted than for ID 1. However, two additional red galaxies are located within  $2''$  in projection toward the north, which could imply imminent additional merging activity.

The core member galaxy with ID 2 is about  $1.5$  mag fainter than the brightest systems, but also shows signs of dynamical assembly activity in the NE-SW directions with another



**Fig. 11.** SED template spectra for an old elliptical galaxy (red), a post-starburst 1 Gyr after SF (green), and an active starburst (blue) redshifted to observed wavelengths at the cluster redshift of XDCP J0044.0-2033 (bottom axis) and in restframe units (top axis). The main filter transmission curves are plotted in black at the bottom of the panel with their names indicated. The corresponding restframe wavelength coverage is shown on top.

NE companion within  $2''$  in projection. Bluer optical emission mainly originates from the SW and NW extensions.

The three spectroscopic members with IDs 4-6 in the outer regions of the cluster have NIR magnitudes in the range  $m^* + 0.5$  to  $m^* + 1$ . Their NIR morphologies are compact and more regular than the core galaxies without any detectable companion within  $2''$ . However, IDs 5 and 6 do exhibit inhomogeneous color distributions within the galaxies.

In summary, the distorted NIR morphologies, color inhomogeneities, and close companions within  $2''$  point toward very active mass assembly of the most central spectroscopic member galaxies, while the secure members in the cluster outskirts exhibit more regular and compact morphologies.

### 3.5. Color–magnitude relation

We now proceed to add quantitative color information and investigate the distribution of the cluster galaxy population in the  $J - K_s$  and  $i - K_s$  versus  $K_s$  color–magnitude diagrams (CMDs). Figure 11 provides a general overview of the observational setup for  $z \approx 1.58$  cluster galaxies by illustrating redshifted template spectra based on Galaxy Evolutionary Synthesis Models (GALEV<sup>2</sup>) (Kotulla et al. 2009) and relevant optical and infrared filter bandpasses (bottom) with their cluster-restframe counterparts (top). The three template spectra depict the expected observed spectral energy distribution (SED) of an elliptical galaxy with an old passively evolving stellar population, a post-starburst template observed one gigayear after the star formation epoch, and the SED of a galaxy with ongoing starburst activity.

At the cluster redshift of XDCP J0044.0-2033, the  $4000 \text{ \AA}$ -break is shifted to the observed  $Y$ -band. The deep HAWK-I observations in  $J$  and  $K_s$  thus probe the SEDs of cluster members

redward of the break, corresponding approximately to a rest-frame  $B - i$  color. With the available deep Subaru imaging data, we obtain an additional measurement blueward of the  $4000 \text{ \AA}$ -break by including the  $i$ -band, which probes the restframe ultra-violet part of the spectrum of cluster galaxies.

The top panel of Fig. 12 shows the NIR HAWK-I  $J - K_s$  versus  $K_s$  color magnitude diagram for galaxies within  $13''$  of the X-ray centroid, at cluster-centric distances of  $13\text{--}30''$  and for all other galaxies at larger distances from the cluster center. The two CMD analysis radii were chosen to provide a 75% statistical cluster member purity for the objects in the very core region based on the overdensity analysis in Sect. 3.1 and a 50% statistical purity for the galaxies. In total, the central  $30''$  region contains about 120 NIR detected galaxies of which about 70 are statistically member galaxies of XDCP J0044.0-2033. The secure spectroscopic members are indicated by cyan (with detected [O II]-emission) and orange squares (without [O II]-emission), while tentative spectroscopic members are marked by smaller black squares.

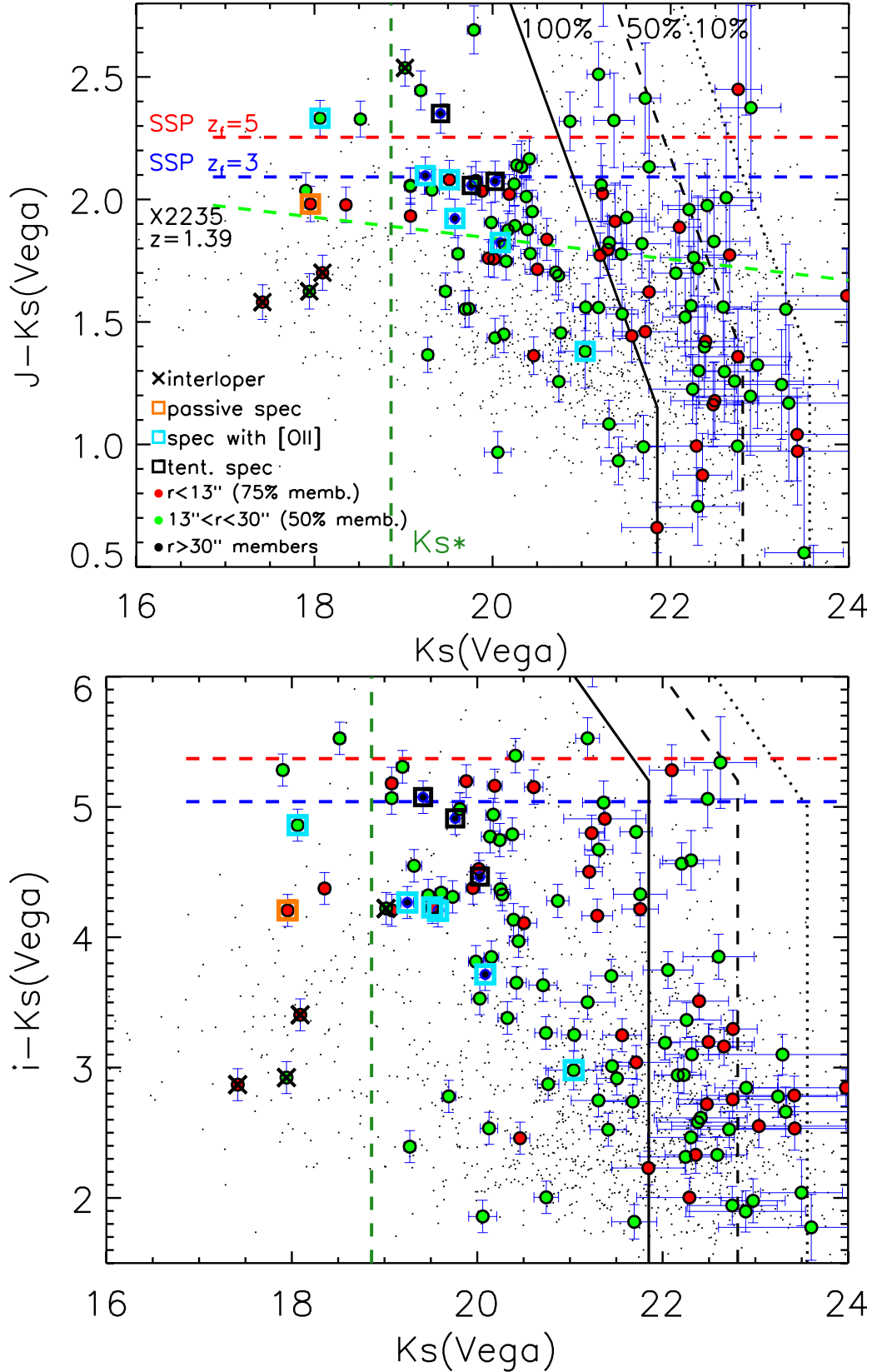
The NIR HAWK-I CMD is based on total  $K_s$  magnitudes and  $1''$ -aperture  $J - K_s$  colors and represents the most accurate photometric data currently available for XDCP J0044.0-2033. Black lines in Fig. 12 indicate the completeness limits of the  $0.5''$ -seeing data as determined in Sect. 2.1 and the vertical green dashed line marks the position of the characteristic SSP model magnitude at the cluster redshift as discussed and observationally confirmed in Sect. 3.3. As visual reference, the horizontal lines depict the predicted colors of solar metallicity SSP models with  $z_f = 3$  and  $z_f = 5$ , while the slanted light green line shows the best fit observed  $J - K_s$  red-sequence location of the cluster XDCP J2235.3-2557 at  $z = 1.396$  found by Lidman et al. (2008).

A clear locus of about 15 galaxies is evident in  $J - K_s$  color-space at the location of the color prediction of the SSP  $z_f = 3$  model in the magnitude range between the characteristic magnitude  $K_s^*$  and  $K_s^* + 1.5$ . This locus would be classically called the cluster red sequence, which has been shown to be in place in massive clusters at least up to  $z \approx 1.4$  (e.g., Strazzullo et al. 2010) and is composed of predominantly passive galaxies with early-type morphologies. However, identifying a well-defined red sequence is not straightforward for XDCP J0044.0-2033 because the spectroscopic members in this locus show indications of active star formation based on their detected [O II]-emission and deformed morphologies due to active merging activity, as discussed in Sect. 3.4. Given the difficulties of separating passive early-type galaxies with high confidence based on the currently limited spectroscopic coverage, we refrain from a quantitative red-sequence fit for the mean color, slope, and scatter at this point. However, the part of the red sequence that appears to be already in place at  $z \approx 1.58$  is fully consistent with the simple stellar-population prediction for a formation redshift  $z_f = 3$  with a negligible slope.

The situation for distinguishing red-sequence galaxies from non-red-sequence objects is even more difficult at the bright end of the galaxy population with magnitudes brighter than the characteristic magnitude  $K_s^*$ . The BCG candidate (ID 1), which is also the only spectroscopic member without a clearly detected [O II]-emission line, is bluer than the discussed locus in  $J - K_s$  color-space and the expectations from lower redshift clusters. The bluer color of this galaxy and the other bright central galaxy (labeled c1) compared with the red-locus population is readily visible in the color-image representation of Fig. 17. The central galaxy region of the spectroscopic member with ID 3a, on the other hand, shows a  $J - K_s$  color that is significantly redder than the red locus, although the spectrum exhibits detected

<sup>2</sup> <http://www.galev.org>





**Fig. 12.**  $J - K_s$  versus  $K_s$  (top) and  $i - K_s$  versus  $K_s$  (bottom) CMDs of the inner region of XDCP J0044.0-2033 for galaxies at cluster-centric distances  $\leq 13''$  (red circles) with a statistical cluster membership of 75% and objects within  $13-30''$  (green circles) with a purity of 50%. Spectroscopic (tentative) members outside  $30''$  are indicated by black circles, small dots mark all other galaxies in the field. Secure (tentative) spectroscopic members are marked by orange/cyan (black) squares with cyan indicating the presence of an [OII]-emission line. The 100%/50%/10% completeness limits are shown by the solid/dashed/dotted black lines. The dashed green line in the top panel depicts the red-sequence location of XMMU J2235.3-2557 at  $z = 1.39$  from Lidman et al. (2008) as reference. Horizontal dashed lines mark solar metallicity SSP-model color predictions for formation redshifts of 3 (blue) and 5 (red), the vertical green line shows the corresponding characteristic  $K_s^*$  magnitude.

[O II]-emission. The region in  $J - K_s$  color-space between the SSP models for  $z_f = 3$  (blue) and  $z_f = 5$  (red), where one would naively expect the bright end of the red-sequence galaxy population to be located, is completely devoid of objects within the  $30''$  analysis region from the X-ray centroid. Based on this observation, we can conclude that the bright end of the red sequence is not yet in place for XDCP J0044.0-2033 for galaxies brighter than the characteristic magnitude  $K_s^*$ .

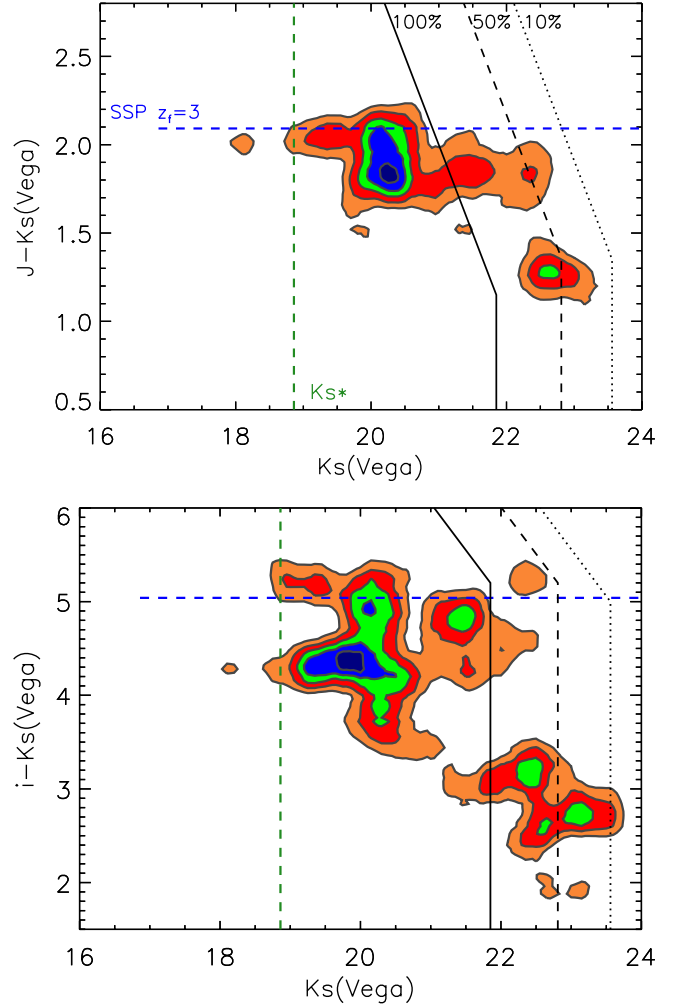
The underlying physical reasons for the significant color deviations of the bright cluster galaxies from simple model expectations are of prime interest within the overall galaxy evolution framework. In particular for galaxies with colors above the red-locus (e.g., ID 3a), the detailed interplay of ongoing star formation activity and possible significant dust attenuation (e.g., Pierini et al. 2004, 2005) remains to be investigated in depth to derive firm conclusions. To this end, we initiated spectroscopic NIR integral-field observations with VLT/KMOS of bright cluster galaxies, which will be discussed in forthcoming publications (Fassbender et al., in prep.).

The situation at the faint end of the red-locus population is equally interesting and is discussed in detail in Sect. 3.6. For now, we examine the bulk of the galaxies within the  $30''$  analysis region with colors significantly bluer than the red-locus population. In these bluer and fainter parts of the CMD the density of the foreground/background populations increases as well, therefore it is not immediately evident in which regions of color-magnitude-space the bulk of the cluster galaxies are located. To overcome this problem and separate the cluster contribution we created a background-subtracted version of the CMD in Fig. 13 (top panel) that illustrates the net cluster galaxy density contours with levels of 3–11 net cluster galaxies per square arcmin per magnitude in  $K_s$  per 0.2 mag in  $J - K_s$  color.

These background-subtracted density contours in the top panel of Fig. 13 show that at least three more components of the cluster galaxy population can be distinguished in addition to the already discussed red-locus population (i): (ii) a “red-sequence transition population” marking the highest density region in the  $J - K_s$  CMD; (iii) a tail extending to fainter  $K_s$  magnitudes with a quasi-constant color offset to the red locus; and (iv) a population of faint and very blue cluster galaxies in the lower right part of the CMD.

The lower panels of Figs. 12 and 13 show the analogous color-magnitude and density diagrams for the  $i - K_s$  color, which now brackets the 4000 Å-break and therefore is more sensitive to ongoing star formation activity. Owing to the 40% increased seeing in the  $i$  band compared to  $K_s$ , the absolute seeing-matched  $i - K_s$  photometric color uncertainties are correspondingly larger than for  $J - K_s$ , which is partially compensated for by the wider spread along the color axis due to the additional additive  $i - J$  contribution to the total  $i - K_s$  color.

As expected, all spectroscopic members with detected [O II]-emission are now shifted down in  $i - K_s$  farther away from the expected color of passive SSP model galaxies. In effect, the locus of red galaxies in  $J - K_s$  is now split into two branches, one in the region of the SSP model expectations, and a second one almost one magnitude bluer. This bluer second branch includes three spectroscopic members with detected [O II] plus the BCG candidate galaxy, as clearly visible in the density plot of Fig. 13 (lower panel). This means that the locus of red galaxies consistent with old passive stellar populations is more sparsely populated than the  $J - K_s$  color, which is less sensitive to residual star formation. The cluster galaxy population of faint blue objects in the lower right corner is again clearly visible in the  $i - K_s$  CMD; it is now distributed over a wider color spread, as



**Fig. 13.** Background-subtracted densities of cluster galaxies of XDCP J0044.0-2033 in the  $J - K_s$  versus  $K_s$  (top) and  $i - K_s$  versus  $K_s$  (bottom) color magnitude diagrams based on the data shown in Fig. 12. The cluster region was constrained to a radius of  $30''$  from the X-ray centroid. The contours and the corresponding filled regions correspond to detected cluster galaxy densities of 3 (orange), 5 (red), 7 (green), 9 (blue), 11 (dark blue) in units of galaxies per arcmin<sup>2</sup> per mag in  $K_s$  per 0.2 mag in  $J - K_s$  for the top panel and densities of 2.3 (orange), 3.6 (red), 4.9 (green), 6.2 (blue), 7.5 (dark blue) galaxies per arcmin<sup>2</sup> per mag in  $K_s$  per 0.4 mag in  $i - K_s$  for the bottom panel. The sharp truncation of the cluster red-sequence at  $K_s \sim 20.5$  and the significant population of faint blue galaxies in the lower right corner are clearly visible. Lines have the same meaning as in Fig. 12.

are the other components identified in the  $J - K_s$ . These different components of the cluster galaxy population are investigated in depth in color-color space in Sects. 3.7 and 3.8, following the examination of the faint end of red-locus galaxies in the next section.

### 3.6. Faint end of the red sequence

To study the faint end of the red sequence, we focus on the  $J - K_s$  versus  $K_s$  CMD to minimize the potential budget of systematics at faint magnitudes due to, for instance, the different instruments, PSF matching, and the sparser population of the red galaxy locus. However, the general conclusion is independent of the color used and is also evident in  $i - K_s$ .

Figure 14 shows the observed histogram of the number of background-subtracted red galaxies as a function of  $K_s$  magnitude for a narrow ( $\pm 0.1$  mag) selected color interval about the SSP  $z_f = 3$  model color (top panel) and a wider ( $\pm 0.25$  mag) color selection (bottom panel). The red lines show the red galaxy counts for the CMD analysis region within  $30''$ , while the blue dashed histogram illustrates the integrated net counts out to the  $R_{500}$  radius. In all cases, independent of color cut and selection aperture, the red galaxy population exhibits a very sharp truncation at  $K_s \approx 20.5$  mag or about  $K_{s*} + 1.6$  after reaching the maximum density in the magnitude interval  $20 \leq K_s \leq 20.5$ . This truncation magnitude of  $K_s \approx 20.5$  mag is about half a magnitude brighter than the 100% completeness limit and even 1.5 mag away from the 50% limit of the HAWK-I data.

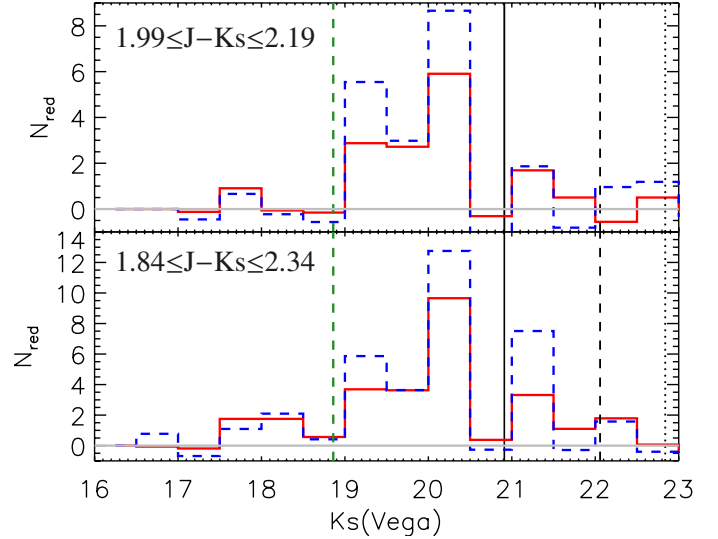
There are several potential systematic effects that might bias the conclusion of a sharp truncation of the red-locus cluster population at  $K_s \geq 20.5$  mag. First, the determined limiting depth of the data might be overestimated so that the apparent suppression at faint magnitudes might be due to incompleteness. That this is not the case was demonstrated with the galaxy number counts in Sect. 2.1. Moreover, this can be visually checked by confirming the clear detections of red-locus galaxies in Fig. 17 (top) at the truncation limit (label “rs1”,  $K_s \approx 20.4$ ) and beyond (label “rs2”,  $K_s \approx 21.2$ ). Second, increasing color errors at faint magnitudes might depopulate the red locus by scattering objects away from the true underlying color. While this is an unavoidable observational bias, we showed in the lower panel of Fig. 14 that a wider color-cut does not change the overall picture. And third, one could invoke that faint red-locus galaxies possibly have larger effective radii than other cluster galaxies and might therefore fall below the surface brightness limit for a detection. However, such a scenario is neither supported by the compact nature of the detected red-locus galaxies (e.g., IDs 4 and 5 in Figs. 10 or 17) nor by other observational high- $z$  studies of passive cluster galaxies (e.g., Strazzullo et al. 2013). In summary, none of the potential systematic biases are likely to influence the apparent lack of red-locus galaxies in the  $K_s$ -magnitude interval of interest  $[20.5, 21.5]$  in a significant way.

We conclude that the population of red-locus galaxies of XDCP J0044.0-2033 is strongly suppressed beyond the surprisingly sharp truncation magnitude of  $K_s \approx 20.5$  mag. This cut-off magnitude in observed red galaxy counts is directly related to the turn-over point of the quiescent galaxy luminosity function, as found previously in a high- $z$  group environment at similar redshift (Rudnick et al. 2012).

### 3.7. Cluster galaxies in color–color space

The availability of deep imaging data in the  $i$ ,  $J$ , and  $K_s$  band allows us to proceed beyond the color–magnitude relation and study cluster galaxies in  $J - K_s$  versus  $i - K_s$  color–color-space to improve the distinction of different object types. Figure 15 shows the color–color space distribution of all detected galaxies with magnitudes  $16 \leq K_s \leq 24$  mag and colored symbols as in the CMDs of Fig. 12, with the additional information of the  $K_s$ -magnitude range indicated by the symbol shape. As visual reference, the colored boxes mark the color–color regions centered on the positions of the template spectra discussed in Fig. 11: red for the passive template, green for the post-starburst galaxy, and blue for the starburst SED.

The lower panel of Fig. 15 illustrates the color–color net cluster density contours after background subtraction. In this representation, the main cluster galaxy population components of the  $30''$  core region become evident as three very significant



**Fig. 14.** Background-subtracted cluster counts of red galaxies in  $J - K_s$  about the color of the SSP  $z_f = 3$  model with a  $J - K_s$  color widths of  $\pm 0.1$  mag (top panel) and  $\pm 0.25$  mag (bottom panel) for the  $R \leq 30''$  core region (red solid lines) and the full projected region within  $R_{500}$  (blue dashed line). The characteristic magnitude  $K_{s*}$  at the cluster redshift is indicated by the green dashed line on the left, the solid/dashed/dotted vertical lines on the right mark the 100%/50%/10% completeness levels for the selected color interval. The sharp cut-off in the red galaxy counts at  $K \approx 20.5$  mag ( $\approx K_{s*} + 1.6$ ) is clearly visible for all histograms.

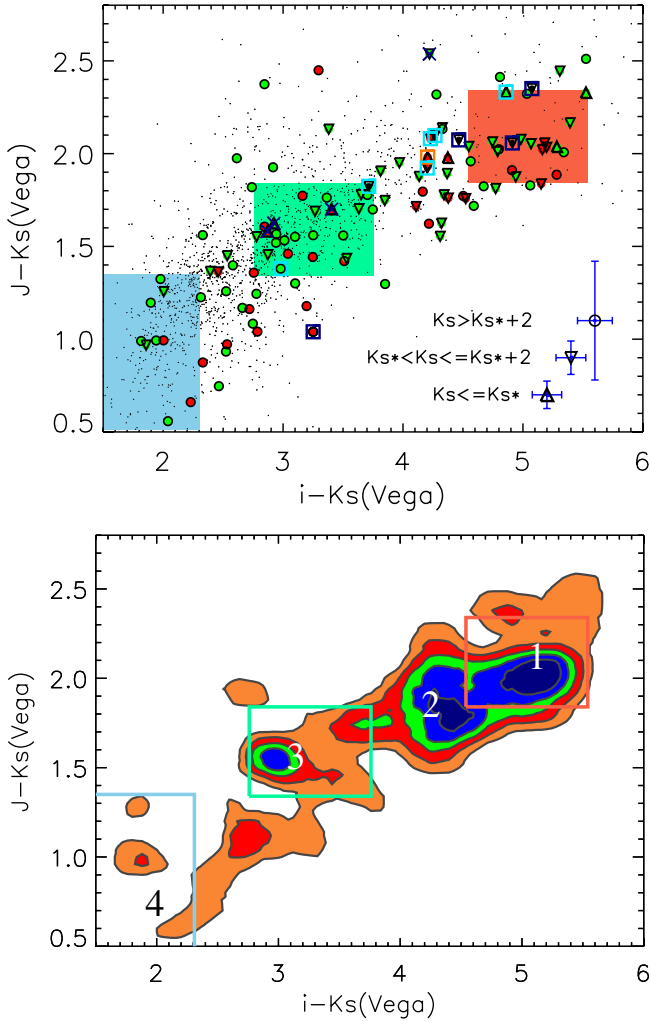
high-density regions plus extensions into the starburst regime in the lower left corner. To more closely investigate the spatial distribution of the different concentrations in color–color space, we defined color–color cuts to distinguish four different galaxy types<sup>3</sup> that represent the four components discussed qualitatively in Sect. 3.5: (1) the red-locus population with colors consistent with a passive SED ( $1.84 \leq J - K_s \leq 2.6$  and  $4.7 \leq i - K_s \leq 5.6$ ); (2) the “red-sequence transition population” ( $1.5 \leq J - K_s \leq 2.2$  and  $3.6 \leq i - K_s \leq 4.6$ ); (3) the post-starburst population ( $1.4 \leq J - K_s \leq 1.8$  and  $2.7 \leq i - K_s < 3.6$ ); and (4) the very blue star-forming/starburst population ( $0 \leq J - K_s < 1.4$  and  $1 \leq i - K_s < 2.9$ ). The centers of these intervals for the color–color selection are marked by the corresponding numbers in the lower panel. For simplicity, we use a general terminology of the color–color selected galaxy classes based on the best-matching SED templates of Fig. 11, which may contain contributions of other galaxy types, for example, dusty starbursts.

### 3.8. Spatial distribution of color-selected galaxies

With the definition of these dominant four subclasses of the distant cluster galaxy population, we can now investigate their spatial distribution within the cluster region and beyond in more detail. The top panels of Fig. 16 show the spatial locations of the four different color–color selected galaxy types, indicated by small colored circles, in the cluster region environment of XDCP J0044.0-2033 with the observed *XMM-Newton* X-ray emission used as gray-scale background image. Logarithmically spaced surface density contours were derived based on the smoothed spatial distribution of objects for each class, which

<sup>3</sup> This is a purely phenomenological classification based on color alone with a naming inspired by the most simple galaxy templates. Contributions of more complex galaxy types (e.g., dust reddened) can be expected, motivating further spectroscopic investigations.





**Fig. 15.** *Top panel:* observed  $J-K_s$  versus  $i-K_s$  color-color diagram of the galaxies shown in Fig. 12 with the same color coding. The  $K_s$  magnitude of the galaxies is indicated by the outer shape: triangles for galaxies brighter than  $K_s^*$  at  $z = 1.58$ , inverted triangles for magnitudes between  $K_s^*$  and  $K_s^* + 2$ , and circles for galaxies fainter than  $K_s^* + 2$ . The median color errors for the three magnitude bins are shown in the legend in the lower right corner. The red, green, and blue boxes show the immediate color-color environment of galaxies with a SED consistent with the template spectra at the cluster redshift (Fig. 11) for a passive galaxy (red), a post-starburst template (green), and active starbursts (blue). *Bottom panel:* background-subtracted cluster galaxy densities in color-color space with density levels of 1 (orange), 2.2 (red), 3.4 (green), 4.6 (blue), 5.8 (dark blue) core cluster galaxies per arcmin<sup>2</sup> per 0.4 mag in  $i-K_s$  per 0.2 mag in  $J-K_s$ . Numbers 1–4 mark the central positions for the color-color selection of different galaxy types with the intervals as specified in the text, while the boxes show the template colors as above: 1: passive, 2: transition objects, 3: post-starburst, 4: starburst/starforming.

are shown with the corresponding colors. The contour levels reflect the full dynamic range of the overdensities above the background levels of the individual classes, which are increasing significantly from the red-locus category (1) with negligible background to the blue starburst category (4) with large foreground/background contributions (see top panel of Fig. 15).

The surface density contours of the red-locus galaxy population span significance levels of 10–120 $\sigma$  above the background and are located mostly within the cluster  $R_{500}$  region, in general correspondence to the extended X-ray emission. The star-forming/starburst galaxy population (cyan, 2.5–10 $\sigma$ )

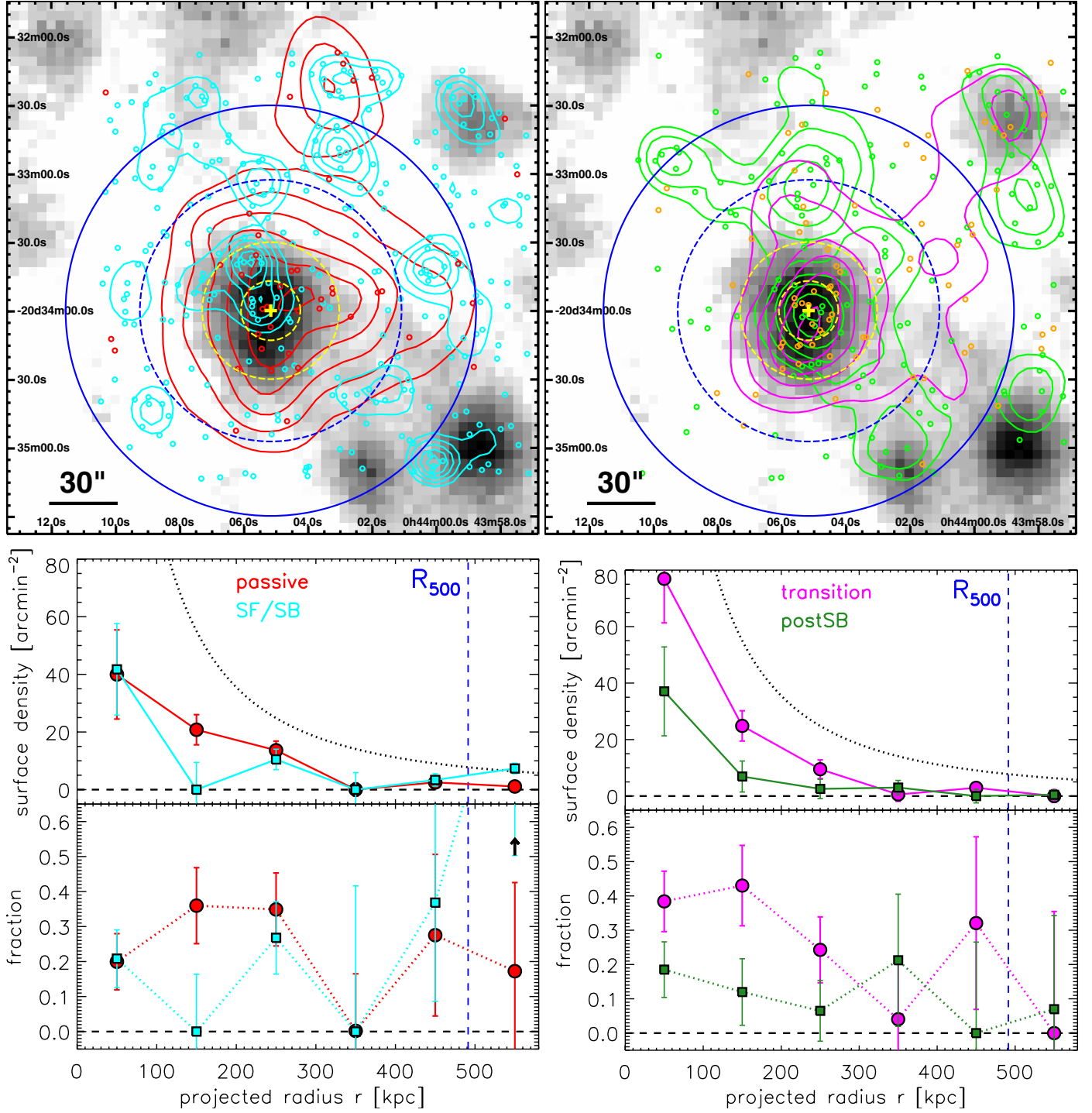
exhibits a 5 $\sigma$  overdensity in the central 100 kpc core region and appears to be mostly concentrated in clumps farther outside, in particular, at the  $R_{500}$  radius and beyond. The categories of red-sequence transition objects (magenta, 4–50 $\sigma$ ) and the post-starburst population (green, 2–15 $\sigma$ ) in the upper right panel exhibit a similar spatial distribution with a compact and sharply peaked core coincident with the regions of maximum X-ray emission.

The lower panels of Fig. 16 show the azimuthally averaged, background-subtracted radial surface-density profiles of the four color-color selected galaxy categories (top) and the galaxy fractions with respect to the total net cluster population (bottom). At radii beyond  $R_{500}$  (rightmost bin) the profiles are consistent with being purely star-forming/starburst dominated within the given significant uncertainties attributed to the outer low-surface density regions. At cluster-centric radii of  $\lesssim 300$  kpc, the environmental effects of the high-density cluster regions becomes evident in the distribution of the overall galaxy population. Within this radius, the surface density of post-starburst systems (category 3), transition objects (2), and red-locus galaxies (1) rise continuously toward the center, most sharply for the red-sequence transition category (magenta), which dominates the cluster population at  $r \lesssim 200$  kpc with a net galaxy fraction of  $\sim 40\%$ . Post-starburst galaxies are found out to 400 kpc with a slowly varying fraction of 10–25%. The fraction of galaxies of the red-locus category (1) reach their maximum of  $\sim 35\%$  in the 100–300 kpc region, before dropping again to  $\sim 20\%$  in the central 100 kpc core region, in stark contrast to low-redshift systems. In the very core region, the contribution of star-forming/starburst galaxies with a fraction of about 20% is similar in number to the red-locus population. The lower panel of Fig. 17 shows a  $V+i+JK_s$  color image of the inner cluster region out to projected distances of  $\approx 300$  kpc with all color-color selected galaxies of the four categories overlaid.

While the azimuthally averaged profiles of the different galaxy types in the lower panels of Fig. 16 are robust in a statistical sense out to the displayed projected radial distance ( $r < 600$  kpc), the individual clumps beyond the  $R_{500}$  radius (dashed circle) suggested by overdensities in the upper panels may not all be part of the surrounding cluster environment. Owing to the increasing foreground/background dominance at these radii, in particular for the bluer galaxy types, some contamination by unrelated foreground structures is to be expected. The suggested main matter-accretion axis onto the cluster coming in from the northern direction as indicated by overdensity contours of all galaxy types, however, is spectroscopically backed up by three secure member galaxies. The currently sparse spectroscopic sampling along the NE to SW axis does not allow any definitive statement on the cluster association of potentially infalling individual group structures beyond  $R_{500}$ . At this point, only the overdensity in the NW (upper right) corner of the FoV featuring weak extended X-ray emission is a spectroscopically confirmed foreground group at  $z_{\text{group}} = 0.721$ , while all other possible substructures have yet to be spectroscopically identified.

### 3.9. Cluster core and formation of the brightest cluster galaxy

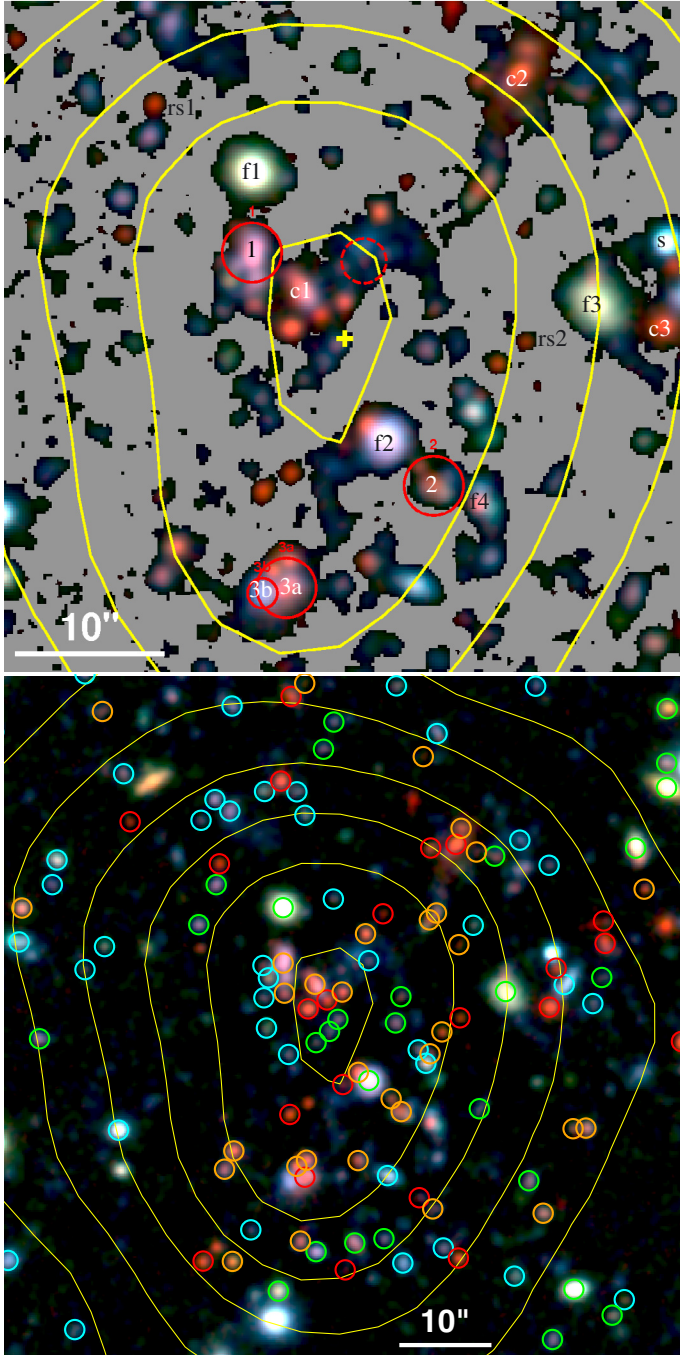
As a final application, we focus on the central 100 kpc core region of XDCP J0044.0-2033 to investigate the current situation and probable future evolution of the formation of the brightest cluster galaxy. As mentioned earlier, we consider the most central secure spectroscopic member with ID 1 (see Table 1 and Fig. 17) as the current BCG. This system is a  $K_s^* - 1$  galaxy



**Fig. 16.** *Top panels:* spatial-excess distributions of color-color-selected galaxies overlaid on the gray-scale *XMM-Newton* X-ray image ( $3.9' \times 3.9'$ ). The blue circles indicate  $R_{200}$  (solid) and  $R_{500}$  (dashed), and the dashed yellow circles mark the  $13''/30''$  analysis radii about the X-ray centroid position (yellow cross). *Top-left panel:* logarithmically spaced density contours (red) of galaxies with colors consistent with a passive SSP SED (red circles) in  $J - K_s$  versus  $i - K_s$  color space and for objects with colors similar to the starburst template (cyan contours and circles). *Top-right panel:* same field showing the spatial distribution of red-sequence transition galaxies (magenta contours and orange circles) and galaxies with post-starburst colors (green contours and circles). *Bottom panels:* corresponding azimuthally averaged and background-subtracted surface-density profiles (upper half) and galaxy fractions (bottom half) as function of projected cluster-centric radius for the same four color-color selected galaxy types as above. The black dotted line indicates the total best-fit galaxy surface-density profile as derived in Fig. 4 as reference. The black upward arrow in the outermost bin of the starburst category (cyan) in the *lower left panel* marks the  $1\sigma$  lower limit for this class.

located at a projected distance from the X-ray centroid of 70 kpc with a low radial restframe velocity offset from the systemic cluster redshift of  $\lesssim 100 \text{ km s}^{-1}$  and no signs of [OII]-emission in its optical spectrum. Because its color is bluer than that of the

red-locus population, it is classified as a red-sequence transition object observed at an epoch of strong merging activity with a number of satellite galaxies, as discussed in Sect. 3.4 and shown in Figs. 10 and 17.



**Fig. 17.** Core view of XDCP J0044.0-2033. *Top panel:*  $45'' \times 45''$  color composite showing the Subaru V- (blue channel) and *i*-band (green) image plus the combined  $J + K_s$  image (red) smoothed with a  $0.5''$  kernel. The black background was rescaled to gray scale for contrast enhancement. The yellow cross indicates the location of the X-ray centroid, whereas contours and symbols for spectroscopic members are as in Fig. 2. The brightest cluster member candidates without available spectroscopic confirmation are labelled “c1-c3”. Obvious low- $z$  foreground galaxies are marked “f1-f3”, “f4” is a spectroscopically confirmed interloper ( $z = 0.592$ ), and a foreground star is indicated by “s”. “rs1” ( $K_s \approx 20.4$ ) and “rs2” ( $K_s \approx 21.2$ ) mark red-locus galaxies at and beyond the found red galaxy truncation magnitude of  $K_s \approx 20.5$ . *Bottom panel:*  $70'' \times 70''$  ( $600 \times 600$  kpc) color composite based on the same data as above. Small colored circles indicate the color-color selected galaxies as defined in Fig. 16.

The cluster-centric offset, merging activity, and fairly blue color are in stark contrast to the more massive dominant central

BCG galaxies in other bona fide distant clusters observed at lower redshifts  $z \approx 1.2$ – $1.4$  (e.g., Collins et al. 2009; Rettura et al. 2010; Strazzullo et al. 2010), that is, at cosmic times 0.5–1.5 Gyr later than XDCP J0044.0-2033. However, this difference in lookback time is longer than the short dynamical friction timescales of  $\lesssim 100$  Myrs for ID 1 and other massive galaxies in the 100 kpc core region as a result of the dynamic interaction with the underlying dark matter halo of the cluster (e.g., Hausman & Ostriker 1978; Binney & Tremaine 1987). This implies that a very significant amount of evolution is expected to occur on these timescales with respect to the BCG mass assembly, structural and color evolution, cluster-centric offset position, and the general reshaping of the densest cluster core region.

The 100 kpc core region already contains an enclosed total stellar mass of about  $M_{*100\text{kpc}} \approx 1.5 \times 10^{12} M_{\odot}$  (Sect. 3.2) and a net cluster excess of about two dozen galaxies (Sect. 3.1) of all color-color-selected galaxy categories (Sect. 3.8). All general ingredients for the build-up of a classical dominant and very massive brightest cluster galaxy are thus already present in the cluster core and are awaiting to be remixed and reshaped within the short dynamical and galaxy evolution timescales in this ultradense region. This process of galactic cannibalism (Hausman & Ostriker 1978) seems to be observed at work here and can be expected to induce a rapid evolution of the BCG properties and its structural appearance.

On timescales of gigayears, the massive cluster galaxies at projected offsets of 150–200 kpc might also end up in the cluster core through dynamical friction and be relevant for the late-time shaping and further evolution of the central brightest galaxy at  $z \lesssim 1$ . This includes the spectroscopic member ID 3a at a current projected offset of 145 kpc and the equally bright, but still spectroscopically unconfirmed, candidate cluster galaxy (labeled c2) located at a projected cluster-centric distance of 180 kpc to the northwest of the core. Both these galaxies are redder than ID 1 and appear to be the central galaxies of an environment rich in satellite galaxies, meaning that they are possibly part of accreted groups. Both galaxies exhibit a total  $K_s$  magnitude within 0.1 mag of the ID 1 luminosity, implying that the current magnitude gap  $\Delta m_{12}$  between the first- and second-ranked galaxy is negligible. To establish the typical observed luminosity gap of  $\Delta m_{12,\text{med}} \approx 0.7$  mag (Smith et al. 2010) between the dominant BCGs and the second-ranked galaxies in low-redshift massive clusters, one could either invoke the discussed expected fast stellar mass growth of the most massive object (ID 1) in the core and/or the coalescence with at least one of the other two bright systems at  $z \lesssim 1$ .

#### 4. Global view of XDCP J0044.0-2033

In this section, we summarize the current knowledge on XDCP J0044.0-2033 with the aim to tie together the results of Sect. 3 and provide a global characterization of the cluster and its immediate large-scale structure environment.

##### 4.1. Dynamical state of XDCP J0044.0-2033

Based on the currently available data from the X-ray to the mid-infrared regime (Santos et al. 2011; Fassbender et al. 2011a, this work), XDCP J0044.0-2033 can be established to be in a quite advanced state of dynamical relaxation, considering that the age of the Universe at the time of observation is only about 4 Gyr. First, the system’s extended X-ray emission detected with *XMM-Newton* (Fig. 2) is compact and to first order symmetrical, with only a slight elongation along the north-south axis.



Second, the redshift distribution of the spectroscopically confirmed member galaxies do not show outliers in velocity space that would indicate dynamical assembly activity along the line of sight. And third, the spatial distribution of the overall galaxy population of XDCP J0044.0-2033 is also compact and highly peaked, with a galaxy density maximum in good spatial concordance with the X-ray centroid position (Sect. 3.1).

Based on this evidence, we can rule out ongoing cluster-scale major-merger activity, although mass-accretion activity with group-scale substructures are still viable at the current resolution of the data. A high degree of cluster virialization implies that the X-ray luminosity-based mass estimate derived from the  $M - L_X$  scaling relation (Reichert et al. 2011) provides a robust total mass estimate, establishing XDCP J0044.0-2033 as a massive bona fide galaxy cluster at  $z = 1.58$ . Consequently, the inferred total stellar fraction (Sect. 3.2) of  $f_{*,500} = (3.3 \pm 1.4)\%$  can also be considered as robust within the specified uncertainties.

Owing to its selection as an extended X-ray source (Fassbender et al. 2011a), that is, independent of its optical/IR properties, XDCP J0044.0-2033 provides an unbiased view of the overall galaxy population of a massive bona fide cluster at a lookback time of 9.5 Gyr. By extrapolating the predicted cluster mass growth histories forward in cosmic time (e.g., Fakhouri et al. 2010), XDCP J0044.0-2033 is expected to be a precursor system of the local very massive cluster population with  $M_{200}(z=0) \sim 2 \times 10^{15} M_\odot$ .

#### 4.2. Global view of the galaxy population in the densest environment at $z = 1.58$

The central  $30''$  (250 kpc) core region of XDCP J0044.0-2033 constitutes the densest galaxy environment currently known at  $z \approx 1.6$ . Visually, this core environment is also richer in galaxies than IDCS J1426.5+3508 at  $z = 1.75$  (Stanford et al. 2012), implying that it is likely the densest galaxy concentration known in the  $z > 1.5$  Universe. This unique  $z \approx 1.6$  cluster galaxy overdensity enabled applying statistical background-subtraction techniques to robustly determine the enclosed stellar mass profile (Sect. 3.2), the  $J$ - and  $K_s$ -band galaxy NIR luminosity functions (Sect. 3.3), background-subtracted color–magnitude relations (Sect. 3.5), and red-galaxy counts (Sect. 3.6), as well as color–color selected spatial distributions and radial profiles of different galaxy types (Sect. 3.8).

The most luminous ( $K_s \lesssim K_{s*}$ ) and most massive cluster galaxies in this ultradense core environment observed at a lookback time of 9.5 Gyr are still found in an active state of development compared with the more evolved galaxies in similarly massive clusters at  $z \lesssim 1.4$ . Evidence for these marked galaxy differences at the observed epoch was seen based on morphological appearance with clear signs of active merging and mass assembly (Sect. 3.4) and their deviating colors with respect to the red-locus population (Sects. 3.5 and 3.7). In particular, the currently brightest cluster galaxy is in a distinct early-evolution stage compared with the most dominant low- $z$  BCGs in terms of total NIR luminosity and therefore stellar mass, cluster-centric offset, and magnitude gap to the second ranked galaxy (Sect. 3.9), but fully consistent with the trends found for high- $z$  cluster samples (e.g., Fassbender et al. 2011a).

At intermediate NIR luminosities with  $K_{s*} \lesssim K_s \lesssim K_{s*} + 2$ , the cluster galaxies in the core span a wide range of colors. A large population of galaxies was found with colors consistent with being in the transition stage between a recent star formation quenching period and their migration phase toward the red sequence (Sect. 3.7). Although a red-locus population

of galaxies with colors expected for old passively evolving systems is present in color–magnitude space at these intermediate luminosities, it is not straightforward to define a classical red sequence of early-type galaxies in XDCP J0044.0-2033 because the red locus still contains a significant fraction of galaxies with detected [O II]-emission lines and/or deformed morphologies (Sect. 3.5). However, we observe a very prominent suppression of red-locus galaxies beyond a magnitude of  $K_s \gtrsim 20.5$ , corresponding to about  $K_{s*} + 1.6$  (Sect. 3.6). This truncation magnitude is remarkably sharp and likely to be associated with the mass-quenching scale at  $z \sim 1.6$  in this dense environment. The observations point toward a preferred  $K_s$  magnitude scale between  $K_{s*} + 1$  and  $K_{s*} + 1.5$  for the upward migration of core cluster galaxies from the blue cloud toward the red sequence in color–magnitude space.

At faint NIR magnitudes ( $K_s \gtrsim K_{s*} + 2$ ) the cluster galaxy population is found to be dominated by blue sources that exhibit colors consistent with star-forming/star-bursting systems. In the very cluster core, this starburst population consists of about 20%. Outside the core, star-bursting cluster galaxies within the  $R_{200}$  volume appear to be clustered in clumps and are not smoothly distributed, especially around the  $R_{500}$  radius and beyond, where this population becomes the predominant galaxy type (Sect. 3.8).

#### 4.3. Environmental effects on the cluster galaxy population

From the spatial distribution and background-subtracted radial profiles for the different color–color selected galaxy types, the influence of the cluster environment in shaping the galaxy population becomes clear (Sect. 3.8). Within projected cluster-centric radii of  $r \lesssim 300$  kpc, the fraction of red-sequence transition objects rapidly increases to  $\sim 40\%$ , accompanied by an increase of the post-starburst fraction to approximately 20%. These two transitional populations dominate the total cluster galaxy budget within the regions of the extended X-ray emission or, correspondingly, at net galaxy surface densities beyond  $\sim 20 \text{ arcmin}^{-2}$ . The star formation quenching process(es) for these transitional galaxies in this high-density region must have acted within about the last 1.5 Gyr since the epoch of observation.

As expected, the fraction of red-locus objects also increases at  $r \lesssim 300$  kpc to a level of  $\sim 35\%$ . However, in the central 100 kpc, this fraction again drops to  $\sim 20\%$ . While the statistical significance of this drop for a single cluster is naturally limited, it might imply that a good fraction of the reddest cluster galaxies are actually accreted as such, for example, as central galaxies of evolved infalling groups. The galaxies with IDs 3a and c2 in Fig. 17 are strong candidates for this latter scenario.

### 5. Discussion

After presenting the results on the galaxy population properties of XDCP J0044.0-2033 in the last two sections, we now compare our findings with previous observational results of high- $z$  cluster environments (Sect. 5.1), with predictions from galaxy evolution models (Sect. 5.2), with field galaxies at similar redshifts (Sect. 5.3), and finally attempt to draw a consistent galaxy evolution picture for the densest high- $z$  environments in Sect. 5.4.

#### 5.1. Comparison with other high- $z$ galaxy clusters

In Sect. 3.3, we measured the total  $J$ - and  $K_s$ -band luminosity functions for the core galaxy population at  $r \lesssim 250$  kpc. Previous studies of the cluster galaxy luminosity function have

consistently found the faint-end slope to be close to  $\alpha \simeq -1$  from low redshifts ( $z < 0.1$ , Lin et al. 2004) to intermediate- $z$  ( $z \lesssim 0.6$ , De Propriis et al. 2013), and up to  $z \lesssim 1.3$  (Strazzullo et al. 2006; Mancone et al. 2012). The three-parameter LF fits for XDCP J0044.0-2033 agree with this fiducial faint-end slope. With a fixed  $\alpha \simeq -1$ , we find characteristic magnitudes  $J^*$  and  $Ks^*$  that fully match the predictions based on a passive luminosity evolution model with a formation redshift of  $z_f = 3$ , which is commonly found in the studies at lower redshifts as well.

However, these findings do not imply that the bright end of the LF is already fully in place at  $z \sim 1.6$ . On the contrary, we showed that the brightest cluster galaxies are only about one magnitude brighter than the characteristic magnitude  $m^*$  and that close-up views reveal clear signs of ongoing merging and mass-assembly activity (Sect. 3.4). Similar conclusions on an active cluster galaxy-mass assembly epoch at  $1.3 < z \lesssim 2$  were reached by Mancone et al. (2010) based on a high- $z$  cluster sample study and for the specific  $z = 1.62$  group environment of XMM-LSS J02182-05102 by Lotz et al. (2013) based on morphological information. Recent claims that ruled out any active mass-assembly in bright cluster galaxies at  $z < 1.8$  by Andreon (2013) evidently do not agree with our direct observations and might be attributable to blending effects in the MIR *Spitzer* data, as shown in Fig. 7.

For the observed evolution of brightest cluster galaxies it has now been established that BCGs grow significantly at  $z \lesssim 1.6$  in NIR luminosity and stellar mass (e.g., Lin et al. 2013; Lidman et al. 2012; Fassbender 2007) driven mainly by mass accretion via merging events (Lidman et al. 2013; Burke & Collins 2013). These results are in line with the direct observations of the major BCG assembly phase of XDCP J0044.0-2033 (Sect. 3.9), which was also reported for a lower mass  $z \sim 1.5$  system by Nastasi et al. (2011). New observational results at the highest accessible cluster redshifts agree qualitatively well with BCG growth predictions from simulations (e.g., De Lucia & Blaizot 2007; Ruszkowski & Springel 2009), which is at odds with previous claims of little evolution in BCGs since  $z \sim 1.5$  (e.g., Collins et al. 2009; Stott et al. 2010).

Significant central star formation activity down to the very cluster core has been reported for most known systems at  $z \gtrsim 1.5$  (e.g., Hilton et al. 2010; Hayashi et al. 2010; Tran et al. 2010; Fassbender et al. 2011c; Tadaki et al. 2012). The inferred fraction of blue starburst/star-forming galaxies of about 20% in the very core of XDCP J0044.0-2033 (Sect. 3.8) therefore is a confirmation of the previously observed trend, although extended to a more massive system in this case. The highest mass cluster with an observed high star formation activity is the recently discovered system SPT-CL J2040-4451 at  $z = 1.48$  with a mass of  $M_{200} \simeq 6 \times 10^{14} M_\odot$  (Bayliss et al. 2013). On the other hand, XMMU J2235.3-2557 at  $z = 1.39$  with  $M_{200} \simeq 7 \times 10^{14} M_\odot$  (Rosati et al. 2009) is the highest redshift cluster with a fully quenched core region within 200 kpc (Strazzullo et al. 2010; Bauer et al. 2011; Grützbauch et al. 2012). Total cluster mass  $M_{200}$  and redshift  $z$  are clearly two key parameters that determine the level of star formation activity in the densest core regions. However, at lookback times of about 9.5 Gyr or more, the assembly histories of the individual clusters must also play an increasingly important role. An early collapse time without major recent mass accretion might then explain the low-mass X-ray group at  $z = 1.61$  with predominantly quiescent galaxies reported by Tanaka et al. (2013) and similar observations for JKCS 041 at  $z = 1.80$  (Andreon et al. 2014; Newman et al. 2014), while most systems at similar redshifts exhibit strong SF activity.

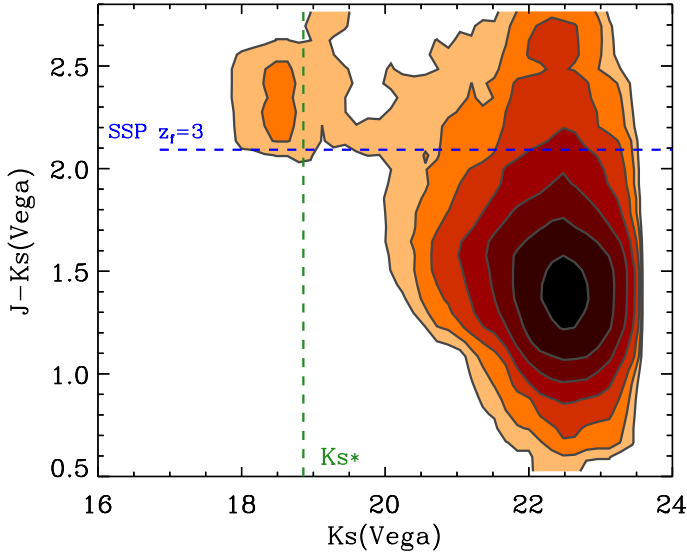
Recent observational studies of the quiescent red-galaxy population and red sequence in distant cluster environments have provided an increasing amount of evidence of significant changes in the old evolved galaxy population in the  $z > 1.4$  regime. Although a significant population of massive, quiescent, early-type galaxies in dense environments has been identified out to  $z \simeq 2$  (Strazzullo et al. 2013; Gobat et al. 2013), the results of Snyder et al. (2012) indicate that most of the stellar mass on the red sequence is assembled at  $z < 2$ . This is in line with the findings by Rudnick et al. (2012), who investigated in detail the quiescent population of the system XMM-LSS J02182-05102 at  $z = 1.62$  and pointed out the importance of merging for the red-sequence build-up. However, the partial red sequences that are already in place in cluster environments at  $z > 1.4$  appear to be all broadly consistent in color and slope with simple expectations from galaxy evolution models where the bulk of stars formed around  $z \sim 3$  (e.g., Snyder et al. 2012; Hilton et al. 2009; Muzzin et al. 2013; Tanaka et al. 2013; Willis et al. 2013). All this agrees well with our results, which showed that there is an active build-up phase of the red sequence of XDCP J0044.0-2033, while at the same time the locus of the existing red galaxy population is consistent with a simple SSP with formation redshift  $z_f = 3$ .

Moreover, the evidence for a sharp truncation of the red galaxy population at magnitudes  $\simeq Ks^* + 1.6$  in Sect. 3.6 confirms the reported trends that the faint end of the red sequence is increasingly underpopulated with increasing redshift (e.g., De Lucia et al. 2004; Koyama et al. 2007; Tanaka et al. 2007, 2013; Stott et al. 2009; Rudnick et al. 2009, 2012; Fassbender et al. 2011b; Lerchster et al. 2011). Assuming a stellar mass of  $\simeq 10^{11} M_\odot$  for a galaxy with characteristic magnitude  $Ks^*$  (Strazzullo et al. 2006), the observed truncation limit in XDCP J0044.0-2033 can be approximated to be on the order of  $M_{*,\text{trunc}} \simeq 0.25 \times M_*(Ks^*) \sim 2.5 \times 10^{10} M_\odot$ . This corresponds well with the red galaxy cut-off scales reported by Tanaka et al. (2013) and Rudnick et al. (2012) for low-mass group environments at similar  $z \simeq 1.6$ .

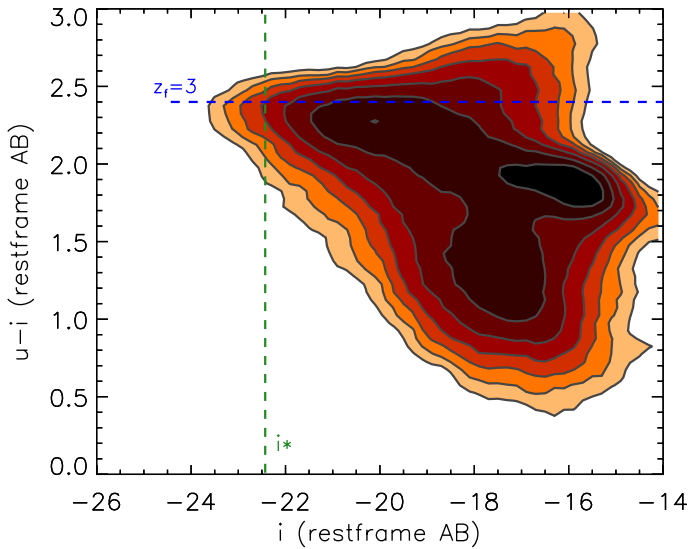
Muzzin et al. (2012) have investigated in detail the environmental effects at  $z \sim 1.0$  based on a sample of rich MIR selected clusters. They found that the population of post-starburst galaxies is increased by a factor of three in cluster environments, based on which, they argued for a rapid environmental-quenching timescale at  $z \sim 1$  or a dominant role of the mass-quenching process at  $z > 1$ . The high fractions of post-starburst and red-sequence transition objects in XDCP J0044.0-2033 (Sects. 3.8 and 4.3), which comprise more than half of the core galaxy population, extend this trend toward higher redshift.

## 5.2. Comparison with model predictions

We now compare predictions based on cosmological simulations and semi-analytic models (SAMs) with our observations of galaxy population properties in the densest  $z \simeq 1.6$  environment. Figure 18 shows the predicted galaxy densities in the  $J - Ks$  versus  $Ks$  CMD based on an extended version of the semi-analytic galaxy formation models by Menci et al. (2008) that also include stellar mass stripping processes and an improved treatment of the metallicity-dependent cooling function. The model CMD is constructed based on the averaged galaxy populations of all objects within halos with a total mass of  $M_{200} \geq 2 \times 10^{14} M_\odot$  in the redshift slice  $1.5 \leq z \leq 1.6$ . The figure is to be compared with the observational analog in Fig. 13 (top panel), with the difference that the rightmost parts of the observed diagram are affected



**Fig. 18.** Predicted densities of galaxies at  $z \approx 1.6$  in the  $J - Ks$  versus  $Ks$  observer-frame CMD for cluster environments with  $M_{200} \geq 2 \times 10^{14} M_{\odot}$  based on the semi-analytic models of Menci et al. (2008). The different contours indicate an increase by factors of two in CMD density from light to dark colors. Dashed lines have the same meaning as in the observed version of the diagram shown in Fig. 13.



**Fig. 19.** Model predictions for the density distribution of cluster galaxies at  $z \approx 1.4$  in the restframe  $u - i$  versus  $i$  CMD based on Bertone et al. (2007). Contour levels increase by factors of 1.5 in CMD density from light to dark colors, dashed lines are the same as in Fig. 13.

by incompleteness and that the observed contours follow linear steps instead of the logarithmic ones of Fig. 18.

The peak in the predicted color–magnitude distribution in Fig. 18 at  $Ks \approx 22.5$  and  $J - Ks \approx 1.3$  exactly coincides with the observed population of low-mass blue star-forming cluster galaxies in Sect. 3.5. After accounting for the approximate factor of two incompleteness corrections in this part of the observed CMD, we conclude that the color–magnitude location of the bulk of the faint blue galaxies agrees with the observations. The model galaxies also agree reasonably well with our results on the bright end of the luminosity function, where the observed magnitudes of the most luminous objects match well, while the predicted colors appear slightly redder than the measured ones.

The most significant discrepancy between the predictions of the model reported by Menci et al. (2008) and the observations occurs in the intermediate-magnitude range  $19 \lesssim Ks \lesssim 20.5$ , where the model is almost devoid of objects owing to fast galaxy evolutionary processes in the model that drive objects toward the bright end on short timescales. For XDCP J0044.0-2033, on the other hand, we find this region to be densely populated by a mix of red-locus galaxies and recently quenched galaxies that transition toward the red sequence. Overall, we conclude that the semi-analytic galaxy evolution model of Menci et al. (2008) agrees well for the starting point of the cluster galaxies in the blue cloud and the end point of massive red galaxies, while the model seems discrepant and does not produce the galaxies found in the transition region in color–magnitude space.

Next, we compare the observations with the implementation of Bertone et al. (2007) of the *Munich galaxy formation model*, which is based on the Millennium Simulation (Springel et al. 2005). Figure 19 shows the restframe  $u - i$  versus  $i$  color–magnitude diagram for all galaxies in cluster environments with  $M_{\text{vir}} \geq 10^{14} h^{-1} M_{\odot}$  at the  $z \approx 1.385$  simulation output snapshot, which was made about half a gigayear later in cosmic time than the observations of Fig. 13. This  $u - i$  restframe color represents the intrinsic galaxy color, which lies in between the observed  $i - Ks$  and  $J - Ks$  measurements, as can be seen in Fig. 11.

The model predictions of Bertone et al. (2007) in Fig. 19 show the bright end of the red sequence already in place. Colors and magnitudes relative to the model lines are fairly consistent with the XDCP J0044.0-2033 observations for the brightest ( $m < m^*$ ) cluster galaxies and, moreover, the highly populated part of red galaxies at intermediate magnitudes ( $m^* \lesssim m \lesssim m^* + 2$ ) agrees well with the observed  $J - Ks$  CMD. Qualitatively, the drop toward bluer colors, that is, the observed red-locus truncation, is also reproduced by the model, although it occurs at fainter magnitudes ( $m \gtrsim m^* + 3.5$ ) in the model. The bulk population of faint blue galaxies at  $m \approx m^* + 4$ , on the other hand, is again fully consistent with the observational findings.

We also checked the predicted  $z \approx 1.6$  cluster galaxy populations based on the  $N$ -body plus hydrodynamical simulations of rezoomed clusters in Romeo et al. (2005, 2008), which include in particular supernova feedback, metal-dependent cooling, and non-instantaneous chemical recycling. For this model, the distant cluster galaxy population shows a distribution in color magnitude space that is in between the two models discussed above.

One common feature for the galaxy evolution models of Bertone et al. (2007), Romeo et al. (2008), and Menci et al. (2008) is a significant population of very red galaxies at faint magnitudes ( $m \gtrsim m^* + 3$ ) in the upper right part of the CMDs. Observationally, such galaxies are notoriously hard to detect because of increasing incompleteness and photometric errors. However, up to the observational limit of the available data this predicted abundant very red, faint cluster galaxy population seems to be suppressed in the measurements.

Overall, the tested galaxy evolution models for distant cluster environments seem to clearly reproduce the observed faint blue bulk population of star-forming galaxies, and the CMD location of the brightest red galaxies is reproduced reasonably well. At intermediate magnitudes, where most of the evolutionary action takes place, the models differ significantly, but seem to commonly underestimate the recently quenched transitioning population of galaxies at  $m^* \lesssim m \lesssim m^* + 1.5$ , as suggested by the presented observations.

Our observed red-galaxy truncation scale of  $M_{*,\text{trunc}} \sim 2.5 \times 10^{10} M_{\odot}$ , however, is consistently predicted by



Gabor & Davé (2012) based on their hydrodynamical cosmological simulation. Here, galaxies passing a critical mass threshold of  $\sim 3 \times 10^{10} M_{\odot}$  enter the mass-quenching regime and are subsequently moving upward in the CMD, reaching the red sequence on a timescale of 1–2 Gyr. Once on the red sequence, the galaxies continue to grow significantly in stellar mass via minor-merging processes, predicting the most massive galaxies to be bluer in color owing to the increasing contribution of infalling low-metallicity, low-mass galaxies. This halo-quenching model thus qualitatively reproduces many of the discussed features of our observed cluster galaxy population in color–magnitude space.

### 5.3. Comparison with field observations

Peng et al. (2010, 2012) have shown that mass-quenching and environment-quenching are distinct and separable processes out to at least  $z \approx 1$ , in the sense that the former acts quasi-independently of the environment and the latter varies little with the halo or stellar mass of the individual galaxy. Previously, Bundy et al. (2006) found evidence for an evolving critical mass-quenching scale above which star formation is suppressed and additional indications that this downsizing process is accelerated in very dense environments, meaning that the mass scale above which feedback processes halt star formation increases toward higher redshifts, while it decreases in denser environments.

In a recent field study at a median redshift of  $z_{\text{med}} \sim 1.8$ , Sommariva et al. (2014) identified a stellar mass threshold for the population of passive galaxies. Here the identified peak of the passive galaxy population as a function of  $K$ s magnitude corresponds well with the identified truncation magnitude of this work, while the sharp decline in number density sets in at a scale about one magnitude fainter. Another finding in common of the field galaxy study by Sommariva et al. (2014) and the observations in this work is the discrepancy between the data and theoretical predictions regarding the presence of red galaxies at faint magnitudes, as was discussed in Sect. 5.2.

### 5.4. Emerging picture of galaxy evolution in the densest $z > 1.5$ cluster environments

After this comprehensive look at the galaxy population properties of XDCP J0044.0-2033, we now attempt to combine all observational results into a more general scenario of the galaxy evolution process in the densest high- $z$  cluster environments. As a starting point, we consider the suggested cluster galaxy evolution toy models by Muzzin et al. (2012) to explain the distribution of the post-starburst population at  $z \sim 1$  and by Rudnick et al. (2012), which describes the build-up of stellar mass on the red sequence since  $z \approx 1.6$ , and extend them as necessary to incorporate the results of this work.

As remarked before, at a lookback time of 9.5 Gyr a number of relevant timescales related to the cluster formation and galaxy evolution process are similar to each other and the lifetime of the cluster itself, which gives rise to the observation of new phenomena in the redshift regime at  $z > 1.5$ . A closer look at these timescales can therefore provide new insights and help in combining many of the recent observational results.

If we assume an early cluster formation epoch for the systems under consideration of  $z \sim 3$ , the cluster lifetime at  $z \approx 1.6$  is at most  $t_{\text{cl,life}} \sim 2$  Gyr, while the accretion and relaxation timescales are  $t_{\text{cl,accr}} \sim t_{\text{cl,relax}} \sim 0.5$ –1 Gyr. Because of the stochastic accretion process of infalling matter and structures

from the surrounding environment, it is immediately evident that we can expect to observe significant cluster-to-cluster variations in the properties of the galaxy and matter content at  $z \approx 1.6$ , depending on the specific formation histories of the individual systems.

For the galaxy population properties in  $z \approx 1.6$  clusters, and in particular, for the formation of the most massive galaxies and the red sequence, we have to consider at least three additional timescales: (i) the environment-quenching timescale  $t_{\text{quench,env}}$ ; (ii) the galaxy mass-quenching timescale  $t_{\text{quench,gal}}$ ; and (iii) the galaxy mass-assembly timescale  $t_{\text{gal,as}}$ , that is, the time needed to double the stellar mass through merging processes.

Galaxies accreted by the cluster from the large-scale structure surroundings are in their majority blue star-forming systems, but can also enter as pre-processed galaxies as part of group environments. Once part of the main cluster halo, the environmental quenching starts to suppress star formation on the timescale  $t_{\text{quench,env}}$  for all galaxies, that is, it causes a parallel shift upward in the CMD. This environmental quenching timescale for XDCP J0044.0-2033 must be longer than the typical residence time of a galaxy in the cluster since accretion (i.e.,  $> 1$  Gyr) to explain the lack of faint red galaxies and the presence of central SF galaxies, and short enough to add a significant number of galaxies to the red sequence in the time span from the cluster formation epoch to  $z \sim 1$  (i.e.,  $\lesssim 4$  Gyr). Based on the observed CMD densities at faint magnitudes and intermediate colors, this timescale for XDCP J0044.0-2033 can be estimated to be on the order of  $t_{\text{quench,env}} \sim 2$ –3 Gyr, which is similar to the cluster age  $t_{\text{cl,life}}$ .

The mass-quenching timescale  $t_{\text{quench,gal}}$ , on the other hand, must be shorter and sets in after galaxies pass the identified stellar mass threshold of  $\log(M_{*,\text{trunc}}/M_{\odot}) \approx 10.4$ , visible as the truncation magnitude of the red-galaxy population in XDCP J0044.0-2033. The mass quenching then drives galaxies with magnitudes brighter than  $Ks^* + 1.6$  upward in the CMD toward the red sequence on a timescale on the order of  $t_{\text{quench,gal}} \sim 1$  Gyr.

The critical third timescale is now the galaxy mass-assembly time  $t_{\text{gal,as}}$ , which is in general longer than the mass-quenching timescale  $t_{\text{quench,gal}}$ , which in most cases results in an evolutionary path along the red sequence through dry-merging activity (e.g., Faber et al. 2007). However, this galaxy mass-assembly timescale strongly depends on the local galaxy density, the relative velocity structure of neighboring galaxies, and in particular, on the local dynamic friction timescale  $t_{\text{DF}}$ . As shown throughout this paper, the central cluster region of XDCP J0044.0-2033 is an extreme environment in these aspects, featuring very high galaxy densities (e.g., Sect. 3.1), a spectroscopic example of low radial velocities for galaxies with IDs 3a and 3b (Sect. 2.2), and a short dynamic-friction timescale in the core (Sect. 3.9). In this favorable environment, the mass-assembly rate is highly accelerated and can become shorter than the mass-quenching timescale for a significant fraction of the cluster galaxies, that is,  $t_{\text{gal,as}} \lesssim t_{\text{quench,gal}}$  or  $t_{\text{gal,as}} \lesssim 1$  Gyr.

This accelerated mass-assembly activity at  $z \approx 1.6$  has now two important consequences for the observed galaxy population properties as objects are driven leftward toward brighter magnitudes in the CMD. First, many galaxies in the cluster core are driven across the mass-quenching threshold  $M_{*,\text{trunc}}$  on a short timescale, where the star formation suppression is accelerated and the rapid transition toward the red sequence begins. Second, as some galaxies continue to rapidly grow in stellar mass through continuing merging activity, the transition toward brighter magnitudes in the CMD can be as fast or faster than the upward

movement, with the observational consequence that a good fraction of the most massive galaxies are still found in their transition state with colors bluer than the expected red sequence, as seen, for example, in Figs. 13 and 15.

This scenario would imply that the immediate influence of the cluster environment at  $z \approx 1.6$  is mainly characterized by the accelerated short galaxy mass-assembly timescale of  $t_{\text{gal,as}} \lesssim 1$  Gyr. As a consequence of this, a large portion of cluster galaxies becomes massive enough within a short time span to pass the mass-quenching threshold and exhibit an increased transitional speed toward the red sequence on the timescale  $t_{\text{quench, mass}} \sim 1$  Gyr. The direct environmental quenching mechanism, on the other hand, must be subdominant and suppresses star formation only on timescales of  $t_{\text{quench, env}} \sim 2\text{--}3$  Gyr. All three relevant timescales that consistently match the observations of XDCP J0044.0-2033 are likely to evolve with decreasing redshift and vary with the total cluster mass  $M_{200}(z)$ .

The proposed scenario, which rests upon a very rapid galaxy mass-assembly timescale  $t_{\text{gal,as}}$  in high- $z$  cluster cores, is fully consistent with Rudnick et al. (2012), who derived a necessary high merger rate at  $z \approx 1.6$  to explain the red-sequence build-up, and Muzzin et al. (2012), who speculated on the dominance of mass quenching over environmental quenching at  $z > 1$ . Moreover, Strazzullo et al. (2013) also tentatively identified the mass-quenching process as the likely dominant mechanism in a  $z \approx 2$  group environment. From the theoretical and modeling side, De Lucia et al. (2012) identified the mass-quenching-dominated regime for galaxies with  $\log(M_*/M_\odot) > 10.5$ , which quantitatively agrees very well with our estimated threshold mass.

In summary, we conclude that in the massive cluster environment of XDCP J0044.0-2033 at  $z \approx 1.58$  we finally directly observe the “fireworks” of massive cluster galaxy assembly at play.

## 6. Summary and conclusions

We have presented a comprehensive study of the galaxy population properties of the massive galaxy cluster XDCP J0044.0-2033 at  $z = 1.58$  based on deep HAWK-I imaging data in the  $J$ - and  $K_s$  bands, complemented by Subaru imaging in  $V$  and  $i$ , *Spitzer* observations at  $4.5\mu\text{m}$ , and new spectroscopic observations with VLT/FORS 2. This X-ray luminous cluster constitutes the densest environment at  $z \approx 1.6$  currently known, which allowed applying statistical background-subtraction techniques to robustly measure and derive various properties of the cluster-associated galaxy population.

Our main results are summarized as follows:

1. Based on the deep HAWK-I near-infrared imaging data, we detected a net excess of about 90 cluster-associated galaxies, of which approximately 75% are concentrated within the central  $30''$  (250 kpc) about the centroid position of the extended X-ray emission.
2. The radial galaxy surface-density profile of the cluster is consistent with a centrally peaked NFW profile with a high concentration parameter of  $c_{200} \approx 10$ . The radial  $4.5\mu\text{m}$  surface brightness profile, on the other hand, is best described by a less compact projected NFW form with  $c_{200} \approx 4.4$ , because the stellar mass peak is offset from the X-ray centroid position.
3. Based on the *Spitzer*  $4.5\mu\text{m}$  imaging data, we measured a total enclosed stellar mass within the projected  $R_{500}$  ( $R_{200}$ ) radii of  $M_{*500} \approx (6.3 \pm 1.6) \times 10^{12} M_\odot$  ( $M_{*200} \approx (9.0 \pm 2.4) \times 10^{12} M_\odot$ ) and a resulting stellar mass fraction of  $f_{*,500} = M_{*,500}/M_{500} = (3.3 \pm 1.4)\%$ , consistent with local values.
4. We measured the total  $J$ - and  $K_s$  band galaxy luminosity functions in the cluster core region within a projected radius of 250 kpc and found characteristic apparent (Vega) magnitudes at the cluster redshift of  $J^* = 20.90 \pm 0.50$  and  $K_s^* = 18.71 \pm 0.66$  for a fixed faint-end slope of  $\alpha = -1$ , in good agreement with predictions from simple stellar population models for a formation redshift of  $z_f = 3$ .
5. Although the overall shape of the galaxy luminosity function seems to be in place at  $z \approx 1.6$ , this does not imply that the bright end of the galaxy population is fully evolved. On the contrary, a closer look at the morphologies and color distributions of the secure spectroscopic members reveals a very active ongoing mass-assembly epoch for the most massive galaxies with clearly visible merger signatures.
6. The  $J - K_s$  and  $i - K_s$  versus  $K_s$  color-magnitude diagrams of the cluster core region revealed a more complex distribution of cluster galaxies in color-magnitude space than lower redshift systems. Although a fairly rich red-locus population is present in the magnitude range  $K_{s*} \lesssim K_s \lesssim K_{s*} + 1.6$  at a color consistent with predictions of  $z_f = 3$  SSP models, a significant number of these galaxies still exhibit [O II] line emission and/or deformed morphologies and thus are not readily identifiable as early-type systems. Moreover, most galaxies at the bright end of the luminosity function at  $K_s \lesssim K_{s*}$  feature colors that are too blue to be associated with a red sequence.
7. The CMDs additionally revealed a large population of galaxies that undergoes color transitioning toward the red sequence, in particular in the magnitude range  $K_{s*} + 1 \lesssim K_s \lesssim K_{s*} + 1.5$ . Even in the cluster core region, we also readily identified the bulk population of faint blue star-forming cluster galaxies at magnitudes around  $K_s \approx K_{s*} + 3.5$ .
8. Number counts along the red-locus galaxy population revealed a sharp truncation magnitude at  $K_s(\text{Vega}) \approx 20.5 \approx K_{s*} + 1.6$ , above which very few red galaxies are present in the cluster up to the detection limit of the data. The faint end of the red-locus population at  $K_s(\text{Vega}) \gtrsim K_{s*} + 1.6$  therefore is not yet fully developed either.
9. We used the  $J - K_s$  versus  $i - K_s$  plane to identify and select different cluster galaxy types in color-color space, distinguishing four significant populations in the cluster core: (1) the red-locus population with colors consistent with a passive SED; (2) a “red-sequence transition population” just bluer of the red locus; (3) a post-starburst population whose colors are consistent with quenched star formation 1 Gyr ago; and (4) the very blue star-forming/starburst population. In the very densest cluster core region at  $r \leq 100$  kpc we identified relative fractions of these galaxy types that comprise about 20% red-locus galaxies (type 1), 40% red-sequence transition objects (type 2), 20% post-starburst galaxies (type 3), and 20% star-forming/starburst systems (type 4).
10. Based on the color-color selection of the four different galaxy types, we also investigated their spatial distribution and radial trends with cluster-centric distance. In terms of absolute number of detected galaxies the first three galaxy types show monotonically rising profiles toward the cluster center, while the star-forming/starburst population appears to be found mostly in clumpy structures, in particular, at and outside the  $R_{500}$  radius, where this blue population becomes

dominant. The relative fractions, on the other hand, indicate a rising fraction of post-starburst galaxies toward the center (type 3), a plateau fraction of 40% for the red-sequence transition galaxies reached within  $r \leq 200$  kpc (type 2), and the red-locus populations that reaches peak fractions of about 35% in the radial range 100–300 kpc and then drops again toward the most central region.

11. We investigated the central cluster core at  $r \leq 100$  kpc with respect to the most likely formation and evolution scenario of the brightest cluster galaxy. The current cluster BCG is identified to be located at a projected cluster-centric distance of 70 kpc from the X-ray centroid with a magnitude of  $K_s^* - 1$ . Its color is blueward of the red-locus population, which places it in the transition-object class (type 2), and it features a highly deformed morphology, which indicates ongoing merging activity along several directions. Based on a measured total stellar mass within  $r \leq 100$  kpc of about  $M_{*100\text{ kpc}} \approx 1.5 \times 10^{12} M_\odot$  and estimated very short dynamical friction timescales of  $\lesssim 100$  Myr, we argued for a very rapid evolution timescale of the current BCG toward increasing stellar masses, redder colors, and smaller cluster-centric offsets, as observed in the majority of cluster cores at lower redshifts.
12. We compared our observational results with model predictions from semi-analytic galaxy evolution models as well as  $N$ -body/hydrodynamical simulations and found that the predictions for the locations of the bulk population of the blue star-forming galaxies (type 4) and the brightest cluster galaxies in color–magnitude space are consistent with the observations, while the predictions for the intermediate transition regions are generally deviant. In particular, we found that the galaxy evolution model of [Menci et al. \(2008\)](#) predicts a transition region that is underpopulated compared with the observations because of a mass evolution toward the bright end that is too fast, while the model of [Bertone et al. \(2007\)](#) predicts the bulk of transitioning galaxies at magnitudes significantly fainter than observed. The qualitatively best match to our results is given by the model predictions of [Gabor & Davé \(2012\)](#), which feature an efficient mass-quenching mechanism above a threshold in stellar mass of  $\sim 3 \times 10^{10} M_\odot$ .
13. Our observational results on the galaxy population properties in the very dense core region of XDCP J0044.0-2033 at  $z \approx 1.6$  point toward a scenario in which environmental quenching is a subdominant process that acts on a timescale on the order of  $t_{\text{quench,env}} \sim 2\text{--}3$  Gyr, while galaxy mass-quenching with a timescale of  $t_{\text{quench,mass}} \sim 1$  Gyr is the more important process to shape the observed galaxy population properties. We see evidence that galaxy evolution in this superdense environment is highly accelerated owing to a rapid mass-assembly timescale through merging processes of  $t_{\text{gal,as}} \lesssim 1$  Gyr, which is for many galaxies similar to or shorter than the mass-quenching timescale. This scenario naturally explains the significant fraction of very bright cluster galaxies that are not yet part of the red sequence, the sharp truncation magnitude of the red-locus population, and the magnitude scale where most galaxies are transitioning toward the red sequence. This would imply that the dominant role of the  $z \approx 1.6$  cluster environment in shaping the observed galaxy population is to drive a large number of core cluster galaxies across the identified mass-quenching threshold of  $\log(M_{*,\text{trunc}}/M_\odot) \approx 10.4$ , while the effect of direct environmental star formation quenching processes are only becoming of increasing relevance at later cosmic epochs.

Owing to its richness in galaxies and the observed ongoing galaxy transformation phenomena, the cluster XDCP J0044.0-2033 at  $z \approx 1.58$  is an excellent testbed to observationally probe in detail galaxy formation processes at the highest matter-density peak at a lookback time of 9.5 Gyr. Upcoming results based on the most recent observations of the cluster galaxies, for example, with VLT/KMOS and the very deep X-ray coverage with *Chandra*, could lead the way to improved insights into the physical mechanisms that shape the observed galaxy population at this epoch.

**Acknowledgements.** The research leading to these results has received funding from the European Union Seventh Framework Programme (FP7/2007-2013) under grant agreement No. 267251 “Astronomy Fellowships in Italy” (AstroFit). This research was supported by the DFG under grant BO 702/16-3. This work is based in part on observations made with the *Spitzer* Space Telescope, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA.

## References

- Andreon, S. 2013, *A&A*, 554, A79
- Andreon, S., Newman, A. B., Trinchieri, G., et al. 2014, *A&A*, 565, A120
- Bauer, A. E., Grützbauch, R., Jørgensen, I., Varela, J., & Bergmann, M. 2011, *MNRAS*, 411, 2009
- Bayliss, M. B., Ashby, M. L. N., Ruel, J., et al. 2013, *ApJ*, submitted [[arXiv:1307.2903](#)]
- Bertin, E., & Arnouts, S. 1996, *A&AS*, 117, 393
- Bertin, E., Mellier, Y., Radovich, M., et al. 2002, in *Astronomical Data Analysis Software and Systems XI*, eds. D. A. Bohlender, D. Durand, & T. H. Handley, ASP Conf. Ser., 281, 228
- Bertone, S., De Lucia, G., & Thomas, P. A. 2007, *MNRAS*, 379, 1143
- Binney, J., & Tremaine, S. 1987, *Galactic dynamics* (Princeton University Press)
- Brodwin, M., Brown, M. J. I., Ashby, M. L. N., et al. 2006, *ApJ*, 651, 791
- Budzynski, J. M., Kopešov, S. E., McCarthy, I. G., McGee, S. L., & Belokurov, V. 2012, *MNRAS*, 423, 104
- Bundy, K., Ellis, R. S., Conselice, C. J., et al. 2006, *ApJ*, 651, 120
- Burke, C., & Collins, C. A. 2013, *MNRAS*, 434, 2856
- Collins, C. A., Stott, J. P., Hilton, M., et al. 2009, *Nature*, 458, 603
- De Lucia, G., & Blaizot, J. 2007, *MNRAS*, 375, 2
- De Lucia, G., Poggianti, B. M., Aragón-Salamanca, A., et al. 2004, *ApJ*, 610, L77
- De Lucia, G., Weinmann, S., Poggianti, B. M., Aragón-Salamanca, A., & Zaritsky, D. 2012, *MNRAS*, 423, 1277
- De Propris, R., Philipps, S., & Bremer, M. N. 2013, *MNRAS*, 434, 3469
- Faber, S. M., Willmer, C. N. A., Wolf, C., et al. 2007, *ApJ*, 665, 265
- Fakhouri, O., Ma, C.-P., & Boylan-Kolchin, M. 2010, *MNRAS*, 406, 2267
- Fassbender, R. 2007, Ph.D. Thesis, Ludwig-Maximilians-Universität München [[arXiv:0806.0861](#)]
- Fassbender, R., Böhringer, H., Nastasi, A., et al. 2011a, *New J. Phys.*, 13, 125014
- Fassbender, R., Böhringer, H., Santos, J. S., et al. 2011b, *A&A*, 527, A78
- Fassbender, R., Nastasi, A., Böhringer, H., et al. 2011c, *A&A*, 527, L10
- Fioc, M., & Rocca-Volmerange, B. 1997, *A&A*, 326, 950
- Gabor, J. M., & Davé, R. 2012, *MNRAS*, 427, 1816
- Galametz, A., Grazian, A., Fontana, A., et al. 2013, *ApJS*, 206, 10
- Gehrels, N. 1986, *ApJ*, 303, 336
- Giodini, S., Pierini, D., Finoguenov, A., et al. 2009, *ApJ*, 703, 982
- Gobat, R., Daddi, E., Onodera, M., et al. 2011, *A&A*, 526, A133
- Gobat, R., Strazzullo, V., Daddi, E., et al. 2013, *ApJ*, 776, 9
- Grützbauch, R., Bauer, A. E., Jørgensen, I., & Varela, J. 2012, *MNRAS*, 423, 3652
- Hausman, M. A., & Ostriker, J. P. 1978, *ApJ*, 224, 320
- Hayashi, M., Kodama, T., Koyama, Y., et al. 2010, *MNRAS*, 402, 1980
- Hilton, M., Stanford, S. A., Stott, J. P., et al. 2009, *ApJ*, 697, 436
- Hilton, M., Lloyd-Davies, E., Stanford, S. A., et al. 2010, *ApJ*, 718, 133
- Kissler-Patig, M., Pirard, J.-F., Casali, M., et al. 2008, *A&A*, 491, 941
- Kotulla, R., Fritze, U., Weilbacher, P., & Anders, P. 2009, *MNRAS*, 396, 462
- Koyama, Y., Kodama, T., Tanaka, M., Shimasaku, K., & Okamura, S. 2007, *MNRAS*, 382, 1719
- Landolt, A. U. 1992, *AJ*, 104, 340
- Lerchster, M., Seitz, S., Brimiouille, F., et al. 2011, *MNRAS*, 411, 2667
- Lidman, C., Rosati, P., Tanaka, M., et al. 2008, *A&A*, 489, 981
- Lidman, C., Suherli, J., Muzzin, A., et al. 2012, *MNRAS*, 427, 550
- Lidman, C., Iacobuta, G., Bauer, A. E., et al. 2013, *MNRAS*, 433, 825



- Lin, Y.-T., Mohr, J. J., & Stanford, S. A. 2004, *ApJ*, 610, 745
  - Lin, Y.-T., Brodwin, M., Gonzalez, A. H., et al. 2013, *ApJ*, 771, 61
  - Łokas, E. L., & Mamon, G. A. 2001, *MNRAS*, 321, 155
  - Lotz, J. M., Papovich, C., Faber, S. M., et al. 2013, *ApJ*, 773, 154
  - Maihara, T., Iwamuro, F., Tanabe, H., et al. 2001, *PASJ*, 53, 25
  - Mancone, C. L., Gonzalez, A. H., Brodwin, M., et al. 2010, *ApJ*, 720, 284
  - Mancone, C. L., Baker, T., Gonzalez, A. H., et al. 2012, *ApJ*, 761, 141
  - Mei, S., Holden, B. P., Blakeslee, J. P., et al. 2009, *ApJ*, 690, 42
  - Menci, N., Rosati, P., Gobat, R., et al. 2008, *ApJ*, 685, 863
  - Muzzin, A., Wilson, G., Yee, H. K. C., et al. 2012, *ApJ*, 746, 188
  - Muzzin, A., Wilson, G., Demarco, R., et al. 2013, *ApJ*, 767, 39
  - Nastasi, A., Fassbender, R., Böhringer, H., et al. 2011, *A&A*, 532, L6
  - Nastasi, A., Scodreggio, M., Fassbender, R., et al. 2013, *A&A*, 550, A9
  - Newman, A. B., Ellis, R. S., Andreon, S., et al. 2014, *ApJ*, 788, 51
  - Ouchi, M., Shimasaku, K., Okamura, S., et al. 2004, *ApJ*, 611, 660
  - Papovich, C., Momcheva, I., Willmer, C. N. A., et al. 2010, *ApJ*, 716, 1503
  - Papovich, C., Bassett, R., Lotz, J. M., et al. 2012, *ApJ*, 750, 93
  - Peng, Y.-J., Lilly, S. J., Kovac, K., et al. 2010, *ApJ*, 721, 193
  - Peng, Y.-J., Lilly, S. J., Renzini, A., & Carollo, M. 2012, *ApJ*, 757, 4
  - Pierini, D., Maraston, C., Bender, R., & Witt, A. N. 2004, *MNRAS*, 347, 1
  - Pierini, D., Maraston, C., Gordon, K. D., & Witt, A. N. 2005, *MNRAS*, 363, 131
  - Reichert, A., Böhringer, H., Fassbender, R., & Mühlegger, M. 2011, *A&A*, 535, A4
  - Rettura, A., Rosati, P., Nonino, M., et al. 2010, *ApJ*, 709, 512
  - Romeo, A. D., Portinari, L., & Sommer-Larsen, J. 2005, *MNRAS*, 361, 983
  - Romeo, A. D., Napolitano, N. R., Covone, G., et al. 2008, *MNRAS*, 389, 13
  - Rosati, P., Tozzi, P., Gobat, R., et al. 2009, *A&A*, 508, 583
  - Rudnick, G., von der Linden, A., Pelló, R., et al. 2009, *ApJ*, 700, 1559
  - Rudnick, G. H., Tran, K.-V., Papovich, C., Momcheva, I., & Willmer, C. 2012, *ApJ*, 755, 14
  - Ruszkowski, M., & Springel, V. 2009, *ApJ*, 696, 1094
  - Santos, J. S., Fassbender, R., Nastasi, A., et al. 2011, *A&A*, 531, L15
  - Schechter, P. 1976, *ApJ*, 203, 297
  - Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, 131, 1163
  - Smith, G. P., Khosroshahi, H. G., Dariush, A., et al. 2010, *MNRAS*, 409, 169
  - Snyder, G. F., Brodwin, M., Mancone, C. M., et al. 2012, *ApJ*, 756, 114
  - Sommariva, V., et al. 2014, *A&A*, submitted
  - Springel, V., White, S. D. M., Jenkins, A., et al. 2005, *Nature*, 435, 629
  - Stanford, S. A., Brodwin, M., Gonzalez, A. H., et al. 2012, *ApJ*, 753, 164
  - Stott, J. P., Pimblett, K. A., Edge, A. C., Smith, G. P., & Wardlow, J. L. 2009, *MNRAS*, 394, 2098
  - Stott, J. P., Collins, C. A., Sahlén, M., et al. 2010, *ApJ*, 718, 23
  - Strazzullo, V., Rosati, P., Stanford, S. A., et al. 2006, *A&A*, 450, 909
  - Strazzullo, V., Rosati, P., Pannella, M., et al. 2010, *A&A*, 524, A17
  - Strazzullo, V., Gobat, R., Daddi, E., et al. 2013, *ApJ*, 772, 118
  - Tadaki, K.-I., Kodama, T., Ota, K., et al. 2012, *MNRAS*, 423, 2617
  - Tanaka, M., Kodama, T., Kajisawa, M., et al. 2007, *MNRAS*, 377, 1206
  - Tanaka, M., Finoguenov, A., Mirkazemi, M., et al. 2013, *PASJ*, 65, 17
  - Tozzi, P., Santos, J. S., Nonino, M., et al. 2013, *A&A*, 551, A45
  - Tozzi, P., et al. 2014, *ApJ*, submitted
  - Tran, K.-V. H., Papovich, C., Saintonge, A., et al. 2010, *ApJ*, 719, L126
  - Willis, J. P., Clerc, N., Bremer, M. N., et al. 2013, *MNRAS*, 430, 134
  - Windhorst, R. A., Cohen, S. H., Hathi, N. P., et al. 2011, *ApJS*, 193, 27
  - Zeimann, G. R., Stanford, S. A., Brodwin, M., et al. 2012, *ApJ*, 756, 115
  - Zhang, Y., Finoguenov, A., Böhringer, H., et al. 2008, *A&A*, 482, 451
- 
- <sup>1</sup> INAF – Osservatorio Astronomico di Roma (OAR), via Frascati 33, 00040 Monteporzio Catone, Italy  
e-mail: rene.fassbender@oa-roma.inaf.it
  - <sup>2</sup> Max-Planck-Institut für extraterrestrische Physik (MPE), Postfach 1312, Giessenbachstr., 85741 Garching, Germany
  - <sup>3</sup> Institut d'Astrophysique Spatiale, Bâtiment 12, Université Paris-Sud, 91405 Orsay, France
  - <sup>4</sup> INAF – Osservatorio Astrofisico di Arcetri, Largo Enrico Fermi 5, 50125 Firenze, Italy
  - <sup>5</sup> European Space Astronomy Centre (ESAC), 7828691 Villanueva de la Canada, Madrid, Spain
  - <sup>6</sup> Australian Astronomical Observatory, PO BOX 915, North Ryde 1670, Australia
  - <sup>7</sup> University of Vienna, Department of Astronomy, Türkenschanzstraße 17, 1180 Vienna, Austria
  - <sup>8</sup> Department of Physics, Durham University, South Road, Durham DH1 3LE, UK
  - <sup>9</sup> National Astronomical Observatory of Japan, Mitaka, 181-8588 Tokyo, Japan
  - <sup>10</sup> Università degli Studi di Ferrara, Via Savonarola 9, 44121 Ferrara, Italy
  - <sup>11</sup> European Southern Observatory (ESO), Karl-Scharzschild-Str. 2, 85748 Garching, Germany
  - <sup>12</sup> Departamento de Astronomía y Astrofísica, Pontificia Universidad Católica de Chile, Casilla 306 Santiago 22, Chile
  - <sup>13</sup> INAF – Osservatorio Astrofisico, via S.Sofia 78, 95123 Catania, Italy
  - <sup>14</sup> CEA Saclay, Orme des Merisiers, 91191 Gif-sur-Yvette, France
  - <sup>15</sup> Centre for Astrophysics and Supercomputing, Swinburne University of Technology, PO Box 218, Hawthorn, VIC 3122, Australia
  - <sup>16</sup> Department of Physics, Ludwig-Maximilians Universität München, Scheinerstr. 1, 81679 Munich, Germany